

Exploring fundamental physics
with gravitational waves

Alfredo Urbano

INFN, sez. di Trieste



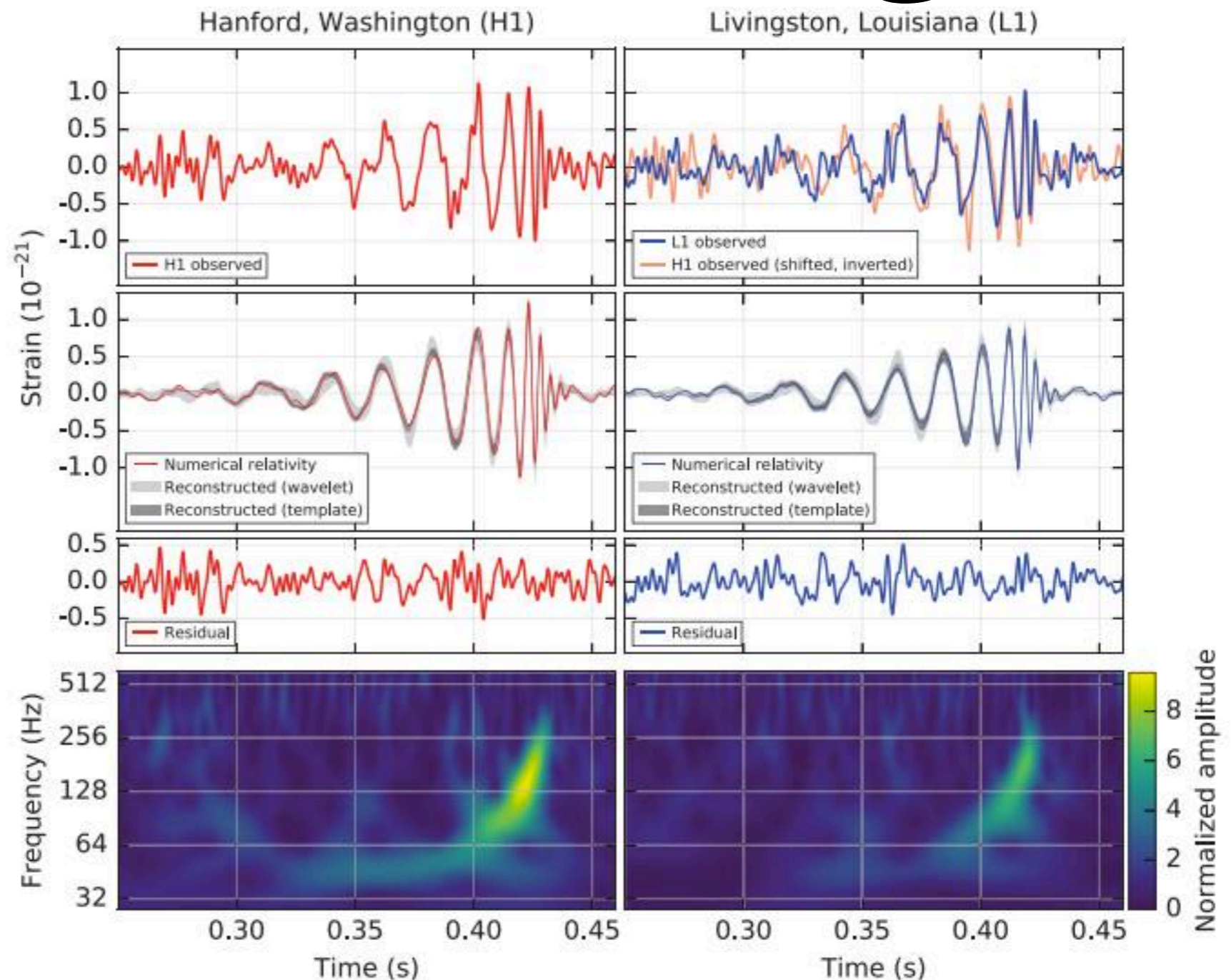
Istituto Nazionale di Fisica Nucleare

Torino, 18 Jan 2019

PART I: Motivations

PART I: Motivations

A new era has begun



PART I: Motivations

Consider a system with characteristic mass M and size L .

$$\Phi \equiv \frac{G_N M}{L}$$

PART I: Motivations

Consider a system with characteristic mass M and size L .

$$\Phi \equiv \frac{G_N M}{L}$$

It is a measure of the Newtonian gravitational potential

$$\phi(r) = - \frac{G_N M}{r}$$

PART I: Motivations

Consider a system with characteristic mass M and size L .

$$\Phi \equiv \frac{G_N M}{L}$$

$$\mathcal{S} = -m \int ds$$

$$ds = \sqrt{g_{\mu\nu} dx^\mu dx^\nu}$$

↓ $g_{\mu\nu} = (-1, 1, 1, 1)$

$$ds = dt \sqrt{1 - v^2}$$

free particle in special relativity

PART I: Motivations

Consider a system with characteristic mass M and size L .

$$\Phi \equiv \frac{G_N M}{L}$$

$$\mathcal{S} = -m \int ds$$

$$ds = \sqrt{g_{\mu\nu} dx^\mu dx^\nu}$$

$$\downarrow F(r) = -m\phi(r)$$

$$g_{00} = 1 + 2\phi(r)$$

$$g_{0i} = 0$$


$$g_{ij} = \delta_{ij}$$

PART I: Motivations

Consider a system with characteristic mass M and size L .

$$ds^2 = - \left(1 - \frac{2G_N M}{r} \right) dt^2 + \frac{dr^2}{\left(1 - \frac{2G_N M}{r} \right)} + r^2 d\Omega^2$$

$$\left[\frac{2G_N M}{r} \right] \sim \Phi$$

$$\frac{1}{\left(1 - \frac{2G_N M}{r} \right)} = 1 + \frac{2G_N M}{r} + \dots$$


PART I: Motivations

Consider a system with characteristic mass M and size L .

$$\Phi_{\text{Sun}} \approx 10^{-6}$$

$$\Phi_{\text{BH}} = 1/2$$

For a BH, $L = 2G_N M$

PART I: Motivations

Consider a system with characteristic mass M and size L .

The parameter Φ is very useful in defining the validity of the Newtonian approximation.

It is not fundamental in characterizing a gravitational field in Einstein's theory.

Indeed, the Einstein field equations are written in terms of the Ricci scalar and the Ricci tensor all of which measure the curvature of the field and not its potential.

PART I: Motivations

Consider a system with characteristic mass M and size L .

$$\mathcal{R}^{1/2} \equiv \sqrt{\frac{G_N M}{L^3}}$$

PART I: Motivations

Consider a system with characteristic mass M and size L .

$$\mathcal{R}^{1/2} \equiv \sqrt{\frac{G_N M}{L^3}}$$

The inverse of the characteristic curvature length scale of the system

PART I: Motivations

Consider a system with characteristic mass M and size L .

$$\mathcal{R}^{1/2} \equiv \sqrt{\frac{G_N M}{L^3}}$$

The inverse of the characteristic curvature length scale of the system

$$R_{\mu\nu} - \frac{1}{2}g_{\mu\nu}R = 8\pi G_N T_{\mu\nu} \quad [T] \sim M/L^3 \quad \Longrightarrow \quad [R] \sim \mathcal{R}$$

PART I: Motivations

Consider a system with characteristic mass M and size L .

$$\mathcal{R}^{1/2} \equiv \sqrt{\frac{G_N M}{L^3}}, \quad [R] = \mathcal{R}$$

example

$$\mathcal{S} = \frac{1}{16\pi G_N} \int d^4x R$$

$\alpha \lesssim 1/\mathcal{R}^{1/2}$: For a given system, we are probing GR down to typical length of order $\alpha \sim 1/\mathcal{R}^{1/2}$

PART I: Motivations

Consider a system with characteristic mass M and size L .

$$\mathcal{R}^{1/2} \equiv \sqrt{\frac{G_N M}{L^3}}, \quad [R] = \mathcal{R}$$

example

$$\mathcal{S} = \frac{1}{16\pi G_N} \int d^4x R (1 + \alpha^2 R)$$

$\alpha \lesssim 1/\mathcal{R}^{1/2}$: For a given system, we are probing GR down to typical length of order $\alpha \sim 1/\mathcal{R}^{1/2}$

PART I: Motivations

Consider a system with characteristic mass M and size L .

$$\mathcal{R}^{1/2} \equiv \sqrt{\frac{G_N M}{L^3}}$$

$$\alpha_{\text{Sun}} \sim 1 / \mathcal{R}_{\text{Sun}}^{1/2} \sim 5 \times 10^8 \text{ km}$$

PART I: Motivations


Consider a system with characteristic mass M and size L .

$$\mathcal{R}^{1/2} \equiv \sqrt{\frac{G_N M}{L^3}}$$

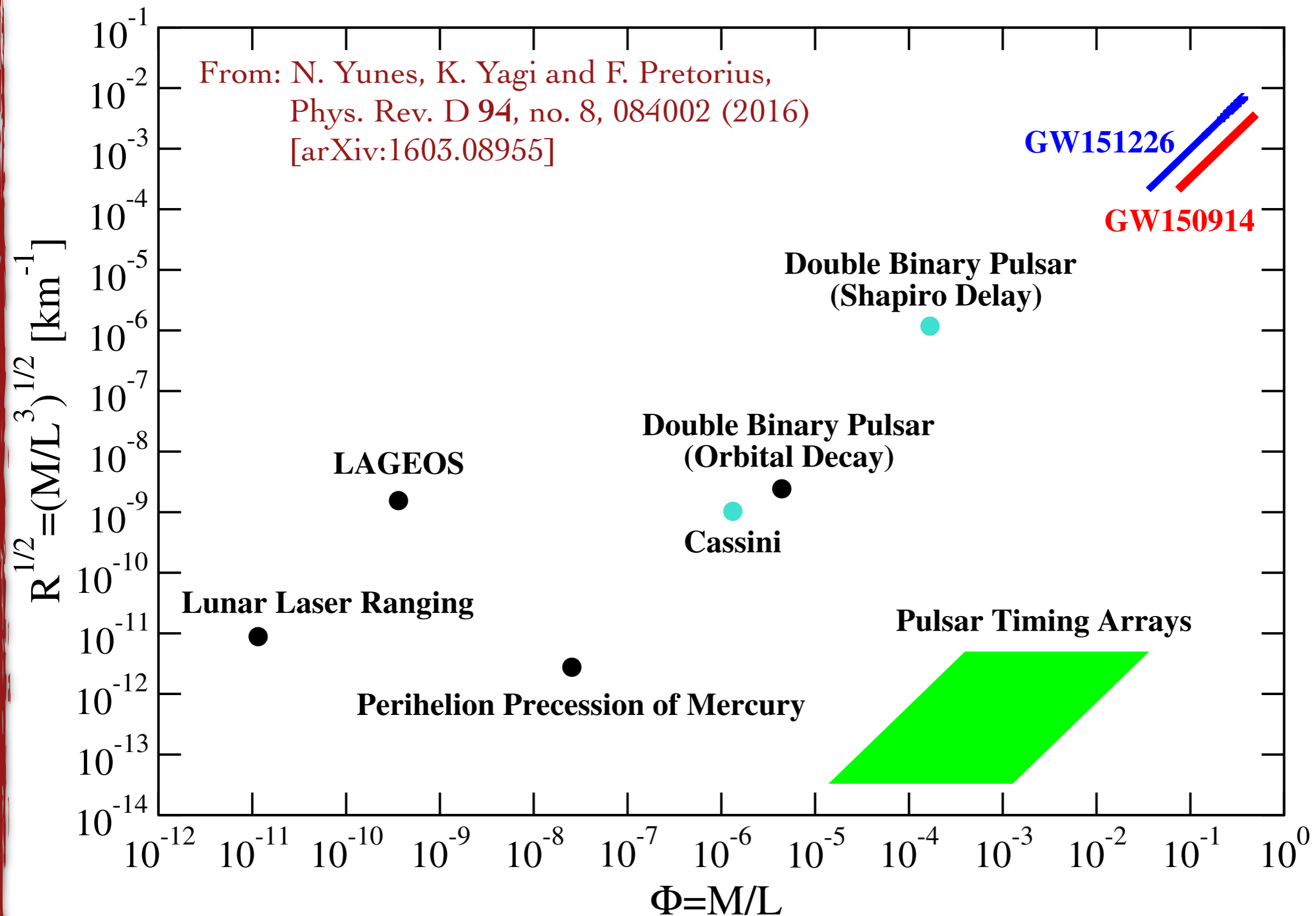
$$\alpha_{\text{BH}} \sim 1 / \mathcal{R}_{\text{BH}}^{1/2} \sim 1 / G_N M$$

$$G_N = 1 / M_{\text{P}}^2$$

$$l_{\text{P}} = 1 / M_{\text{P}}$$

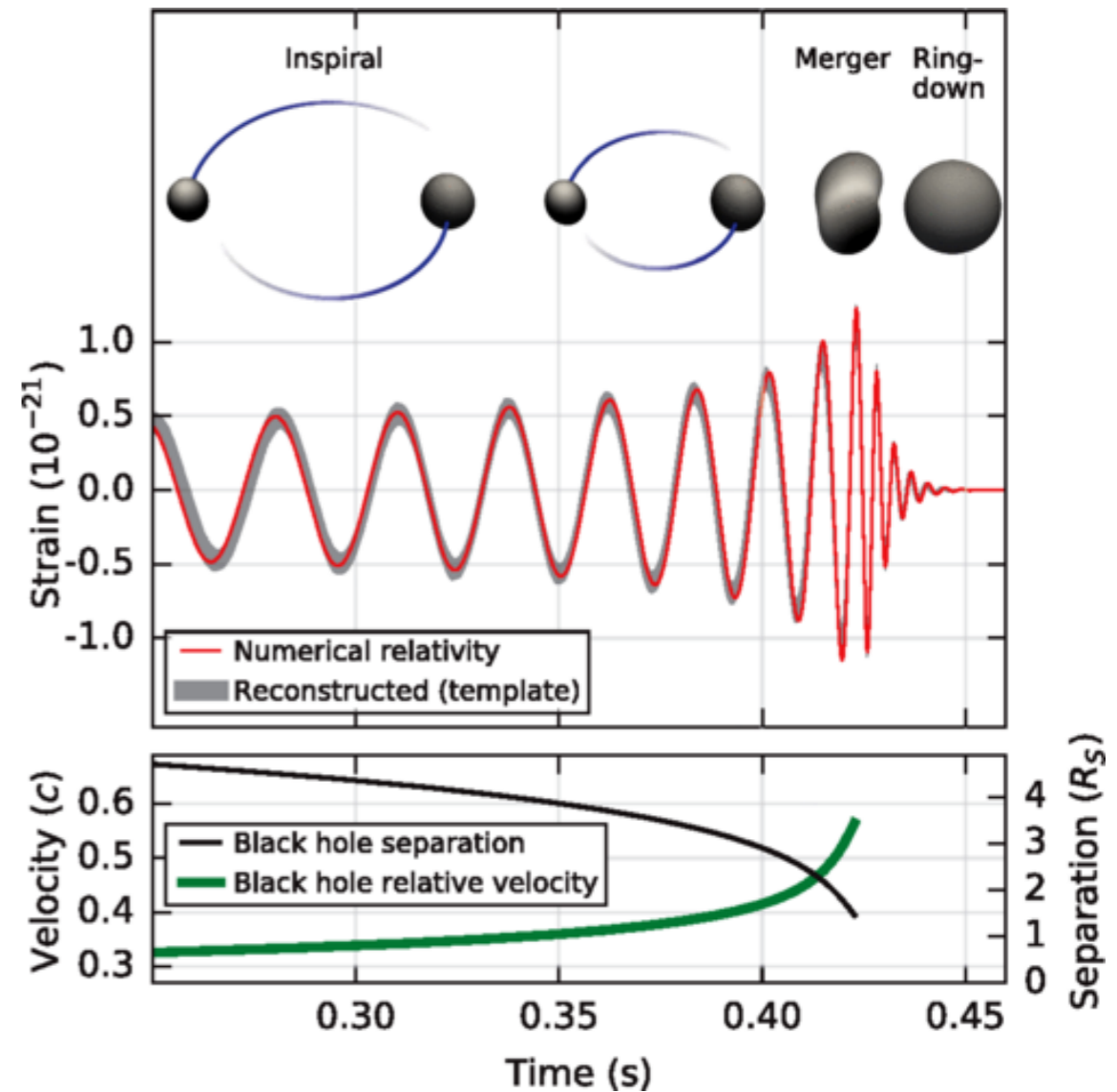

$$\alpha_{\text{BH}} \sim l_{\text{P}} \left(\frac{M}{M_{\text{P}}} \right)$$

PART I: Motivations



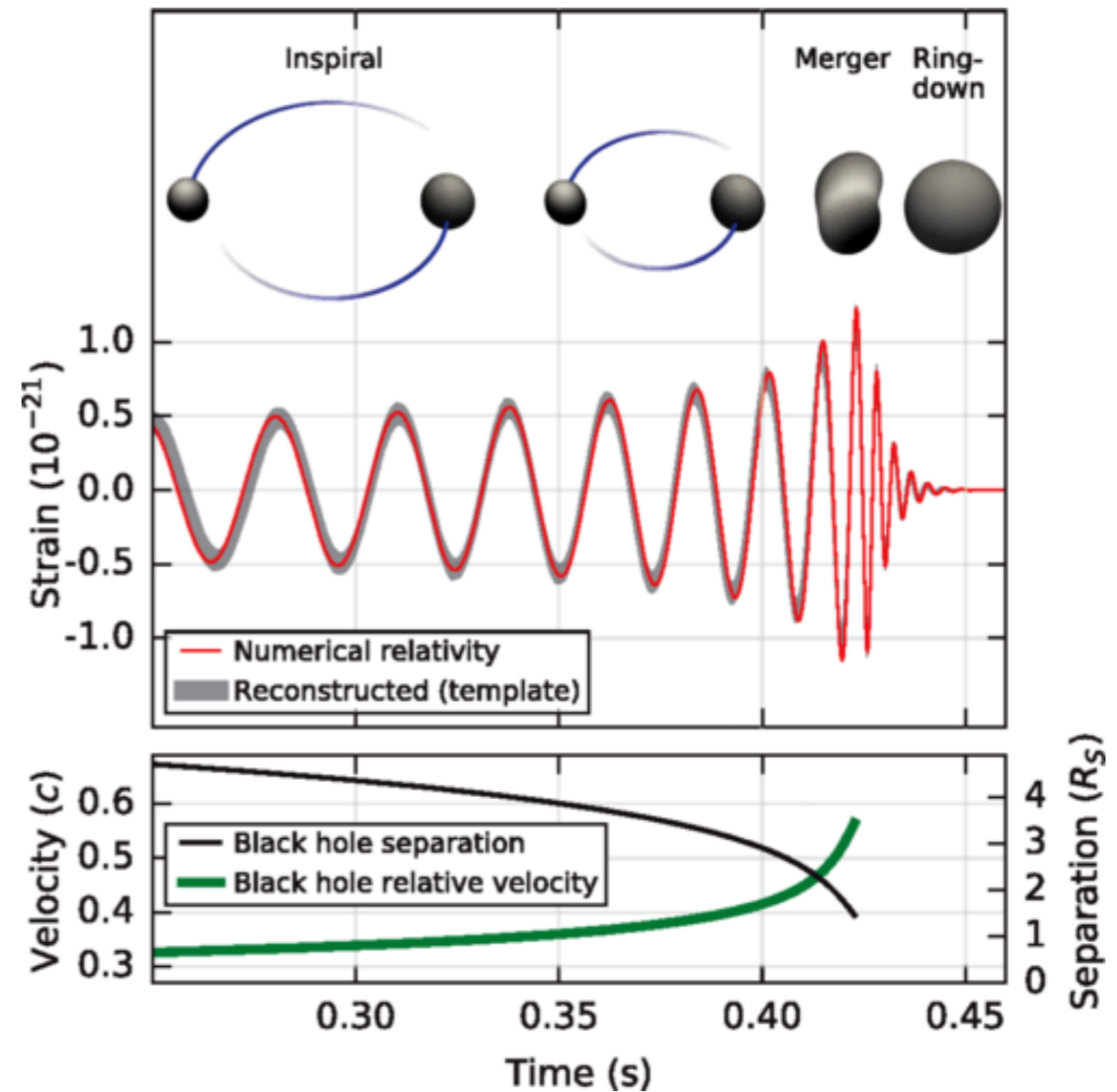
PART I: Motivations

Inspiral phase: The compact objects are well separated with respect to the total mass.



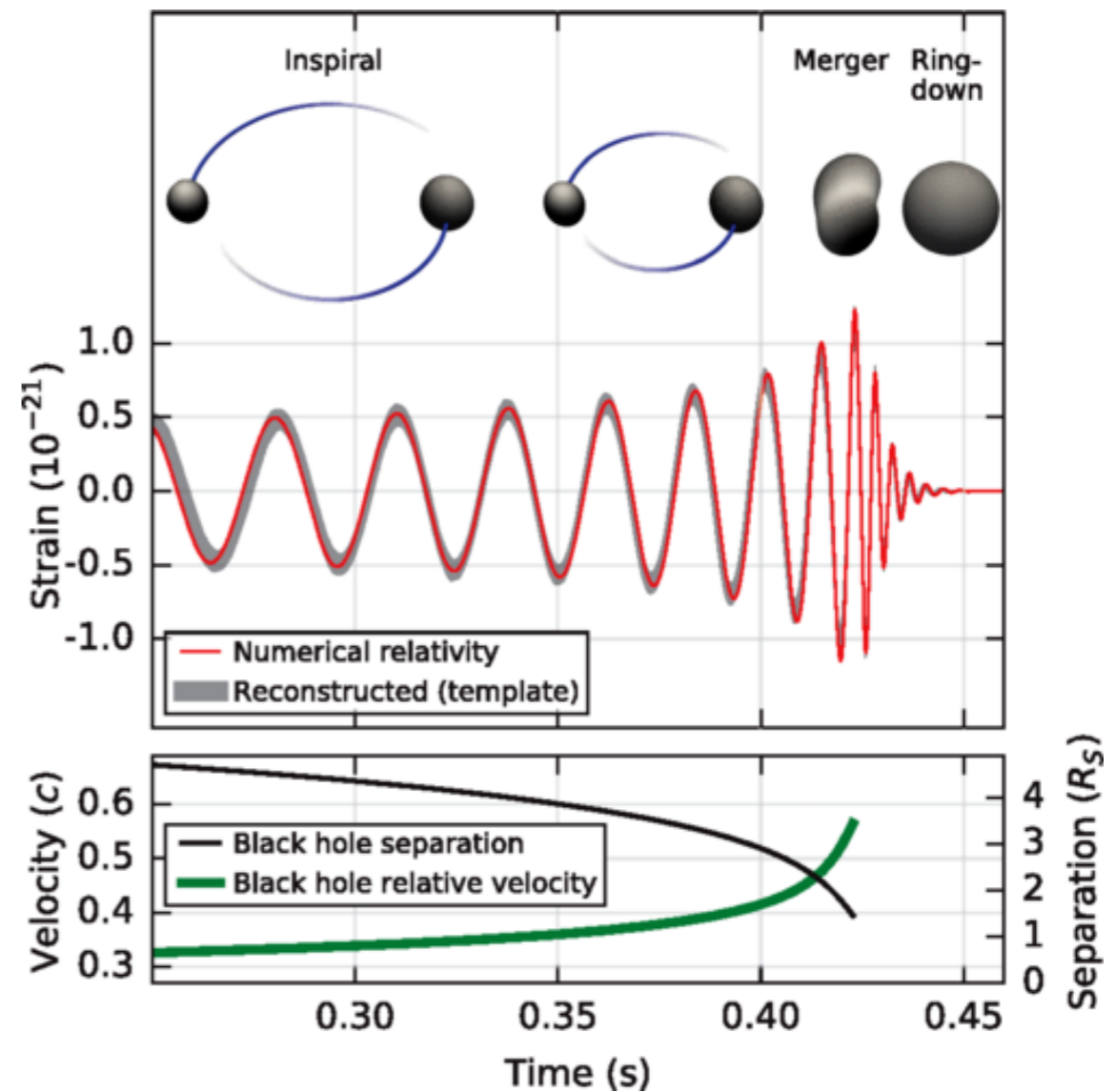
PART I: Motivations

Plunge & merger:
The compact objects are very close to each other and velocities approach the speed of light. The two objects coalesce.



PART I: Motivations

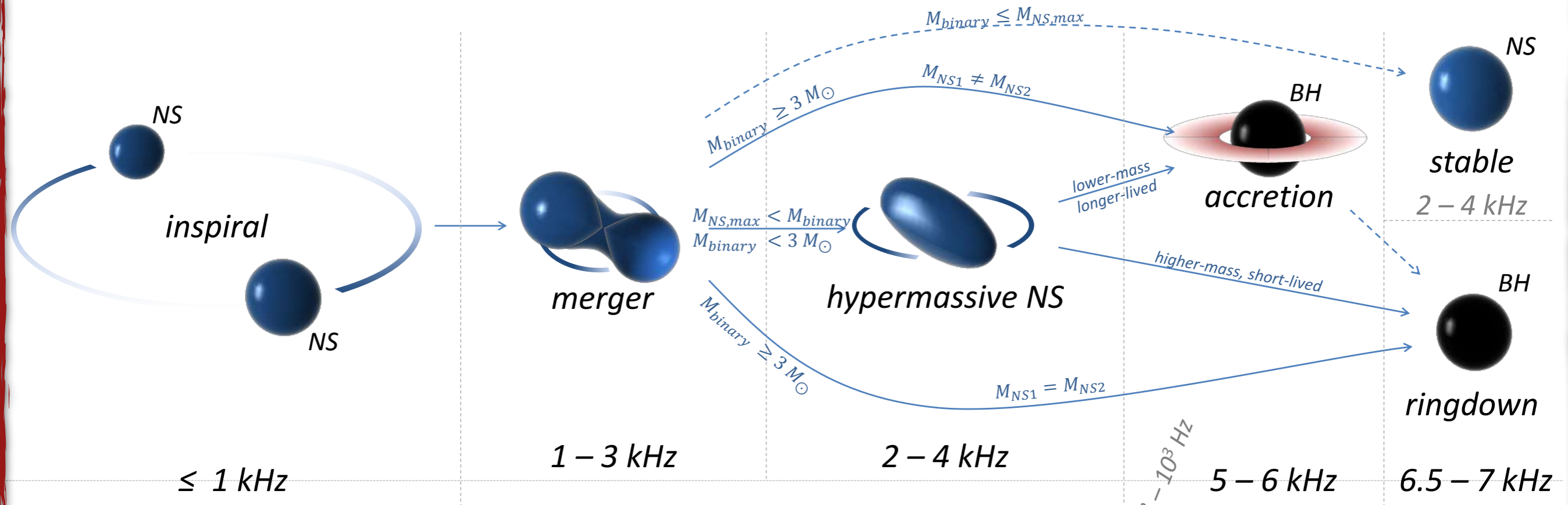
Ringdown: The highly distorted remnant formed after merger oscillates, radiating away any deformations and relaxes to a stationary state.



This classification is clean in concept, and applies to both BH-BH and NS-NS merger.

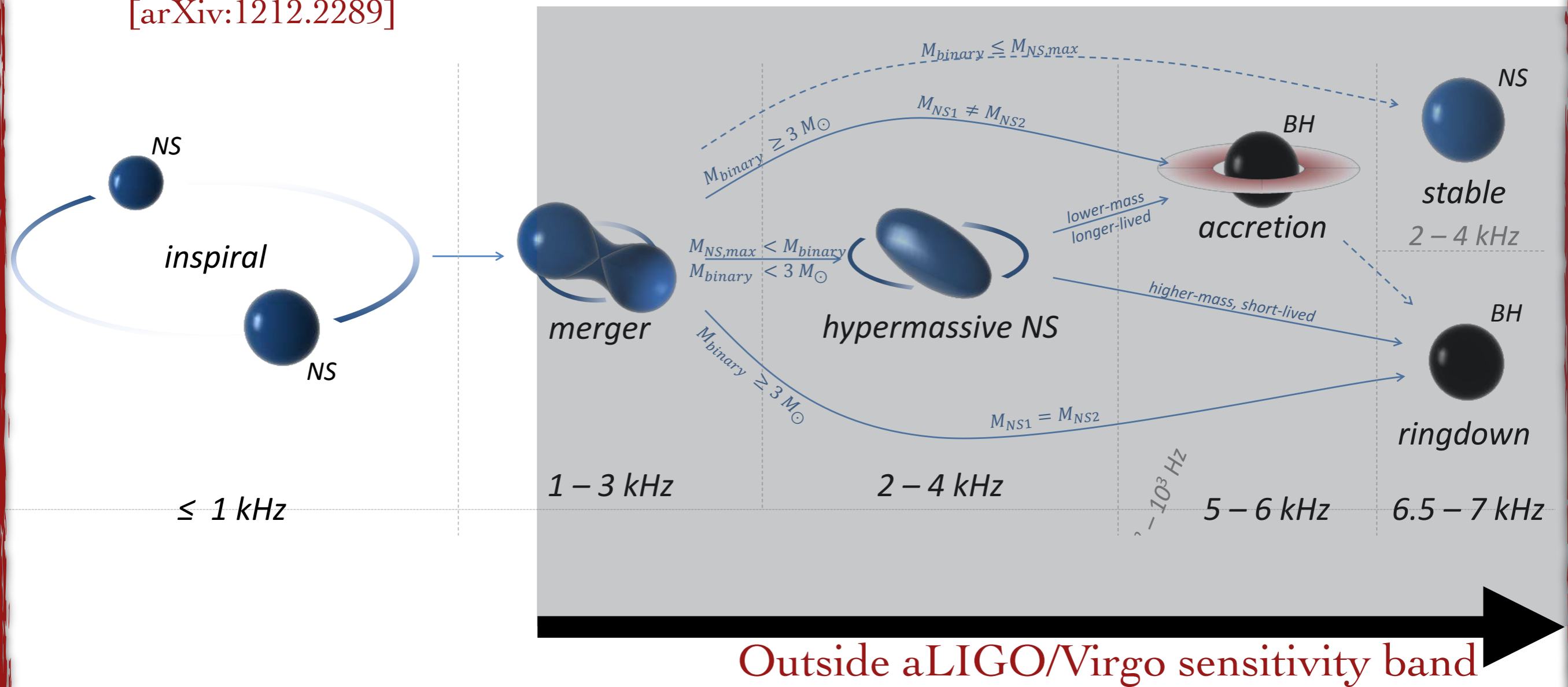
PART I: Motivations

From: I. Bartos, P. Brady and S. Marka,
 Class. Quant. Grav. **30**, 123001 (2013)
 [arXiv:1212.2289]

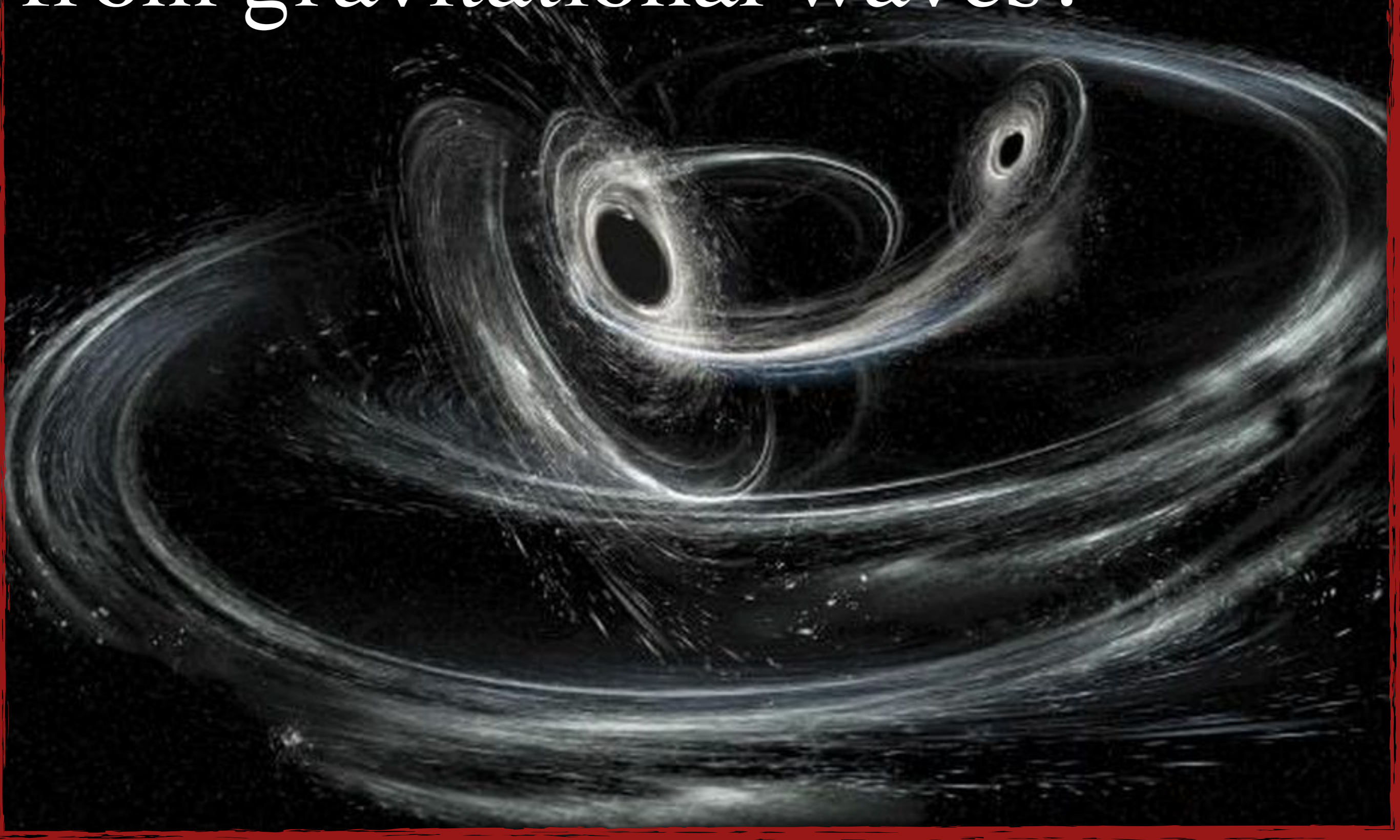


PART I: Motivations

From: I. Bartos, P. Brady and S. Marka,
Class. Quant. Grav. **30**, 123001 (2013)
[arXiv:1212.2289]



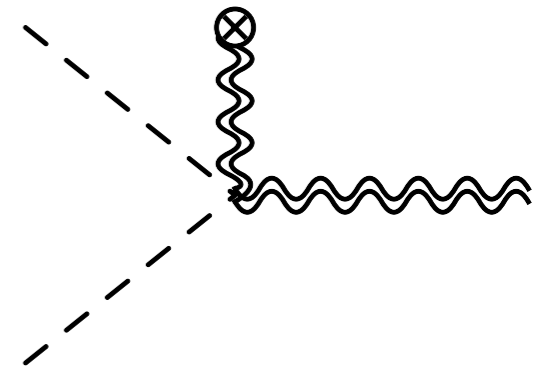
What we can learn
from gravitational waves?



Black hole as “particle detector”

GW as a probe for new light and weakly-coupled (dark) particles

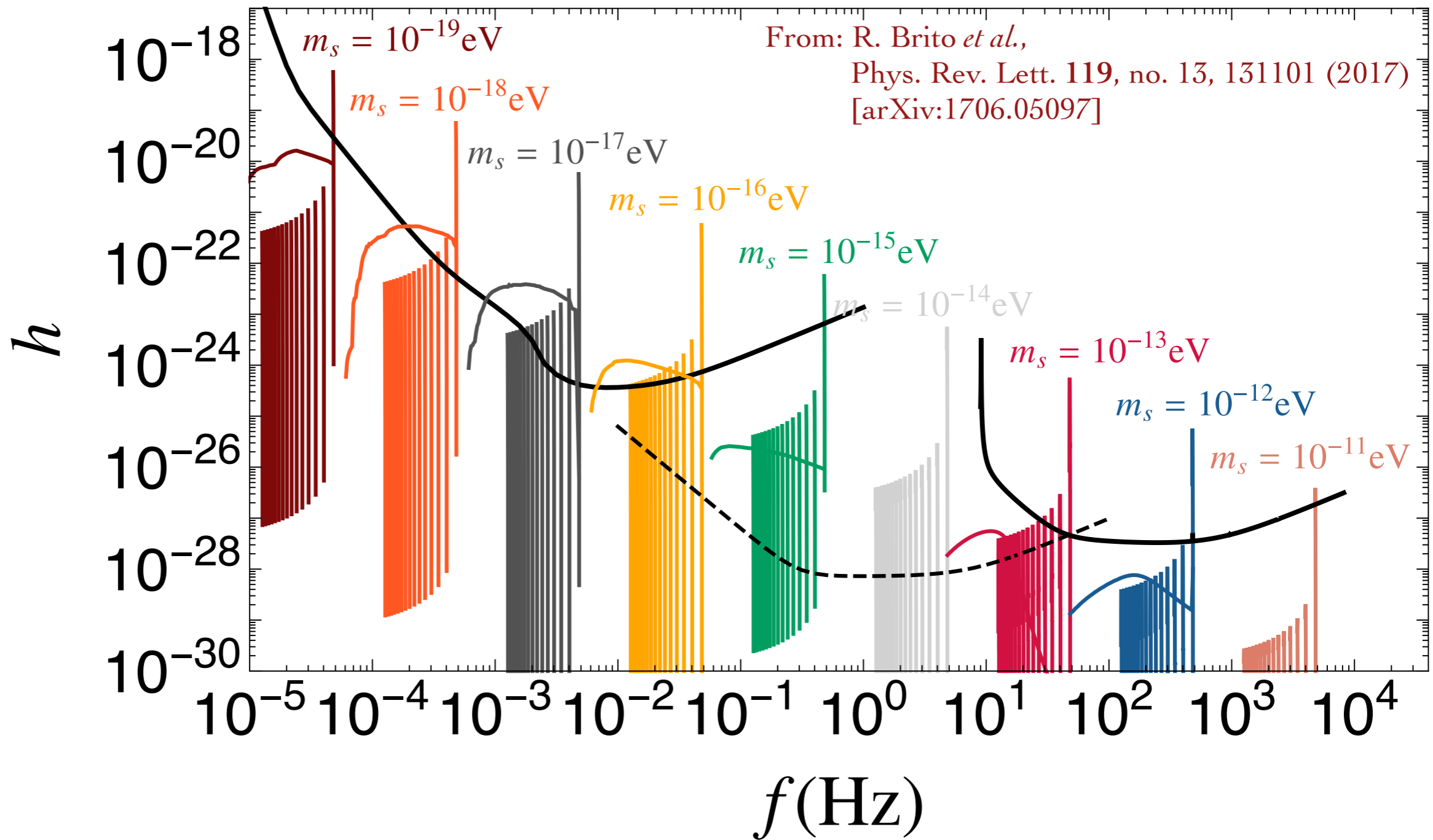
The condensate is dissipated through the emission of GWs (with frequency set by the scalar field mass)



$$M \approx 6.7 \left(\frac{10^{-12} \text{ eV}}{m} \right) M_{\odot}$$

$$\frac{m}{10^{-12} \text{ eV}} \simeq 240 \text{ Hz}$$

Black hole as “particle detector”

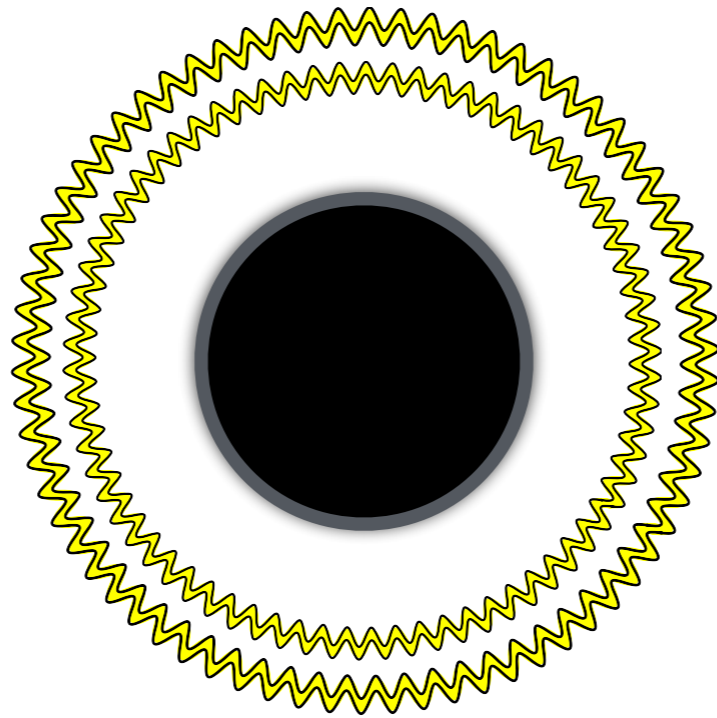


$$M \approx 6.7 \left(\frac{10^{-12} \text{ eV}}{m} \right) M_{\odot}$$

$$\frac{m}{10^{-12} \text{ eV}} \simeq 240 \text{ Hz}$$

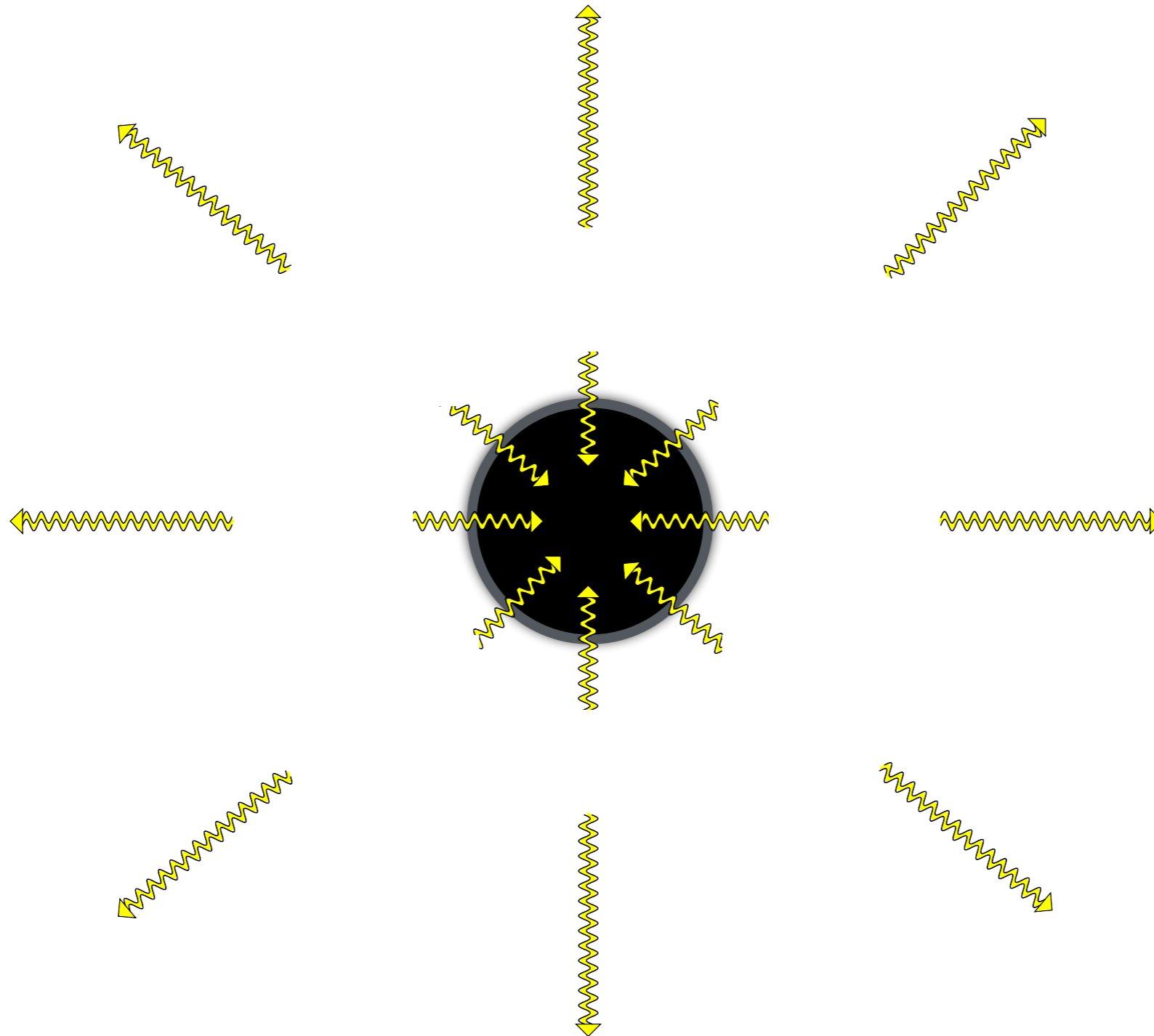
Gravitational wave “echoes”

GW as a probe for deformation of the black hole event horizon compared to GR predictions



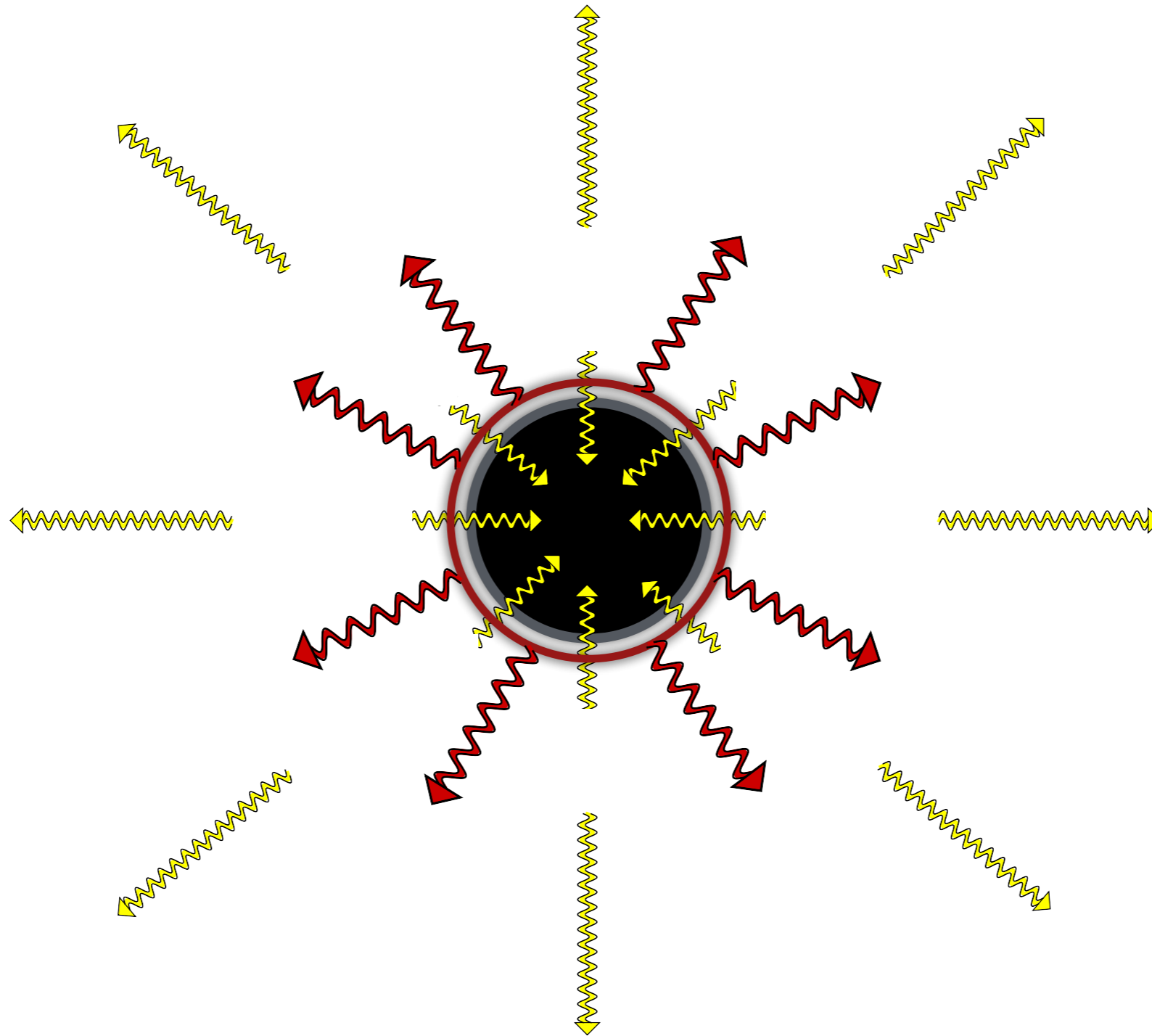
Gravitational wave “echoes”

GW as a probe for deformation of the black hole event horizon compared to GR predictions



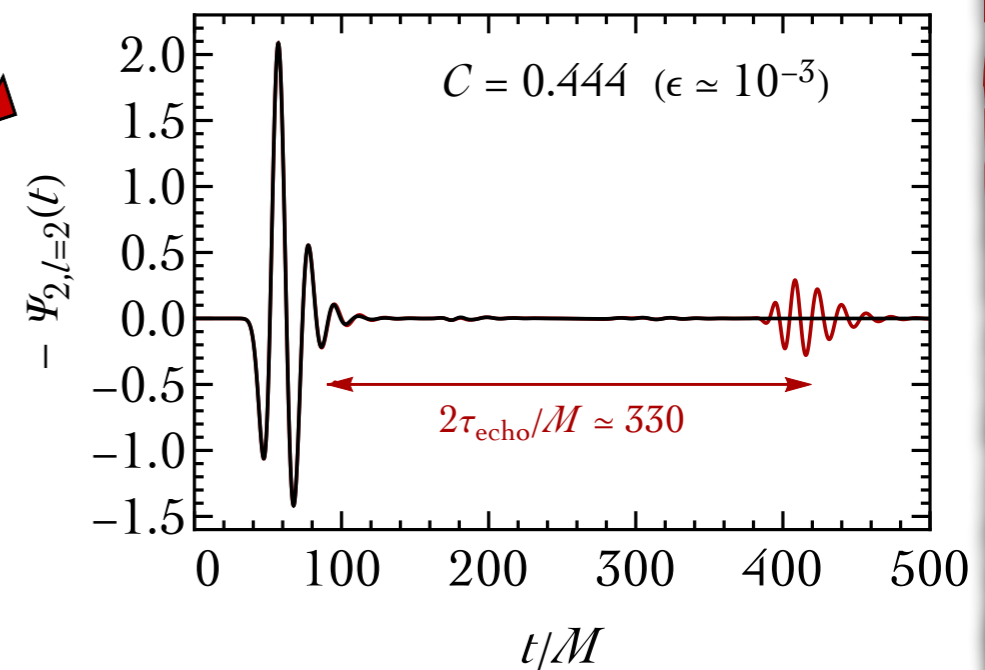
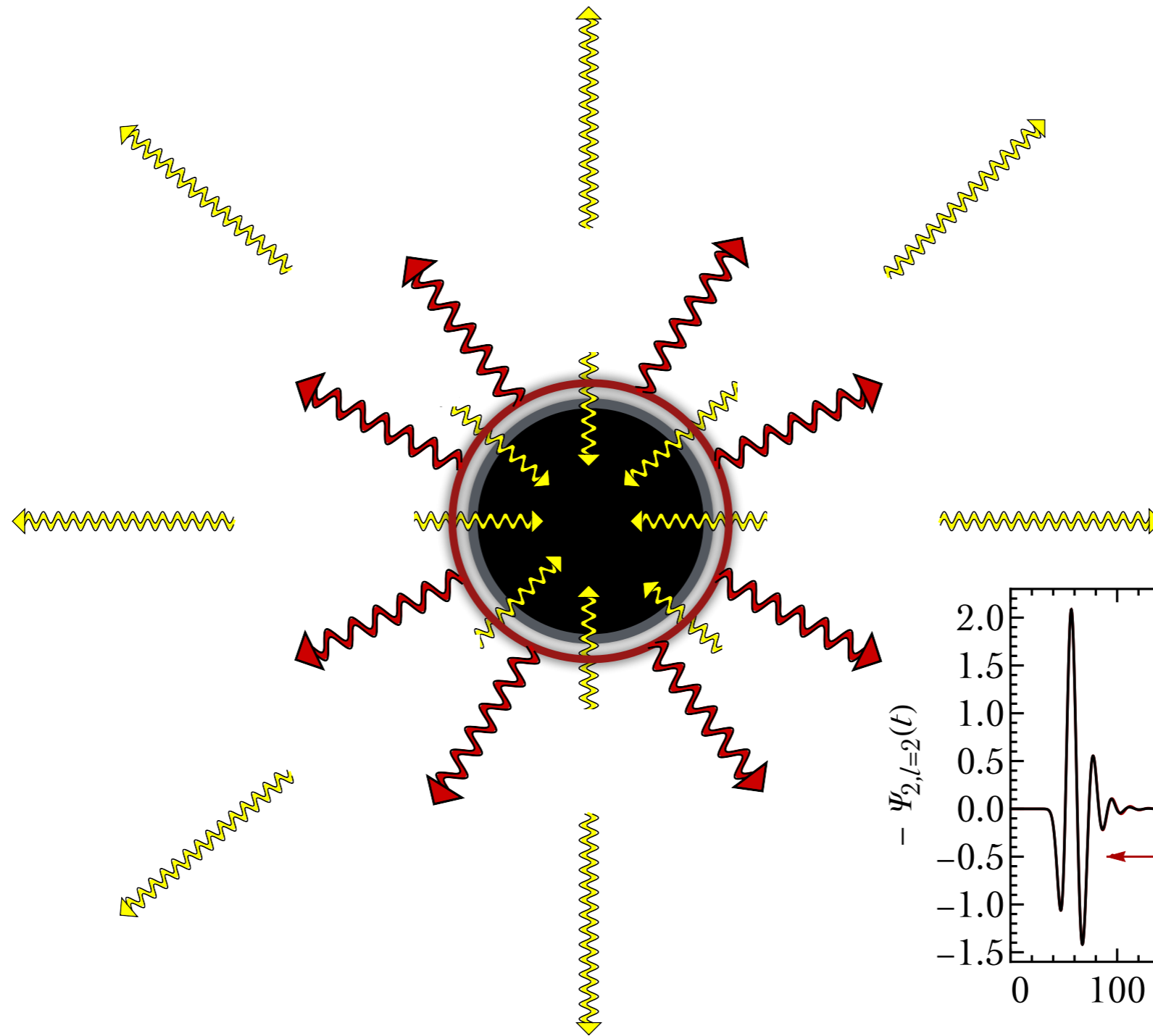
Gravitational wave “echoes”

GW as a probe for deformation of the black hole event horizon compared to GR predictions



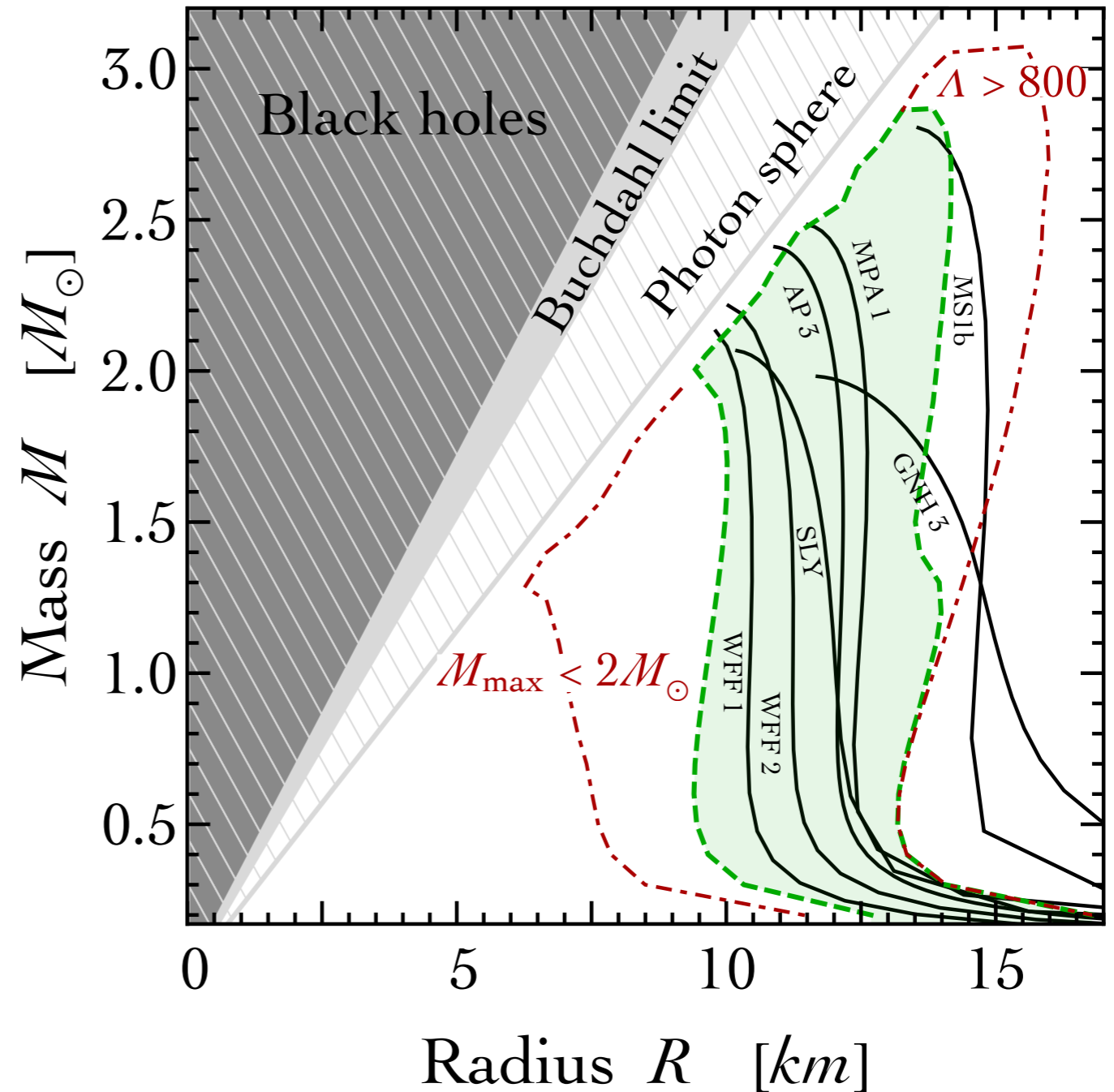
Gravitational wave “echoes”

GW as a probe for deformation of the black hole event horizon compared to GR predictions



Neutron star EoS

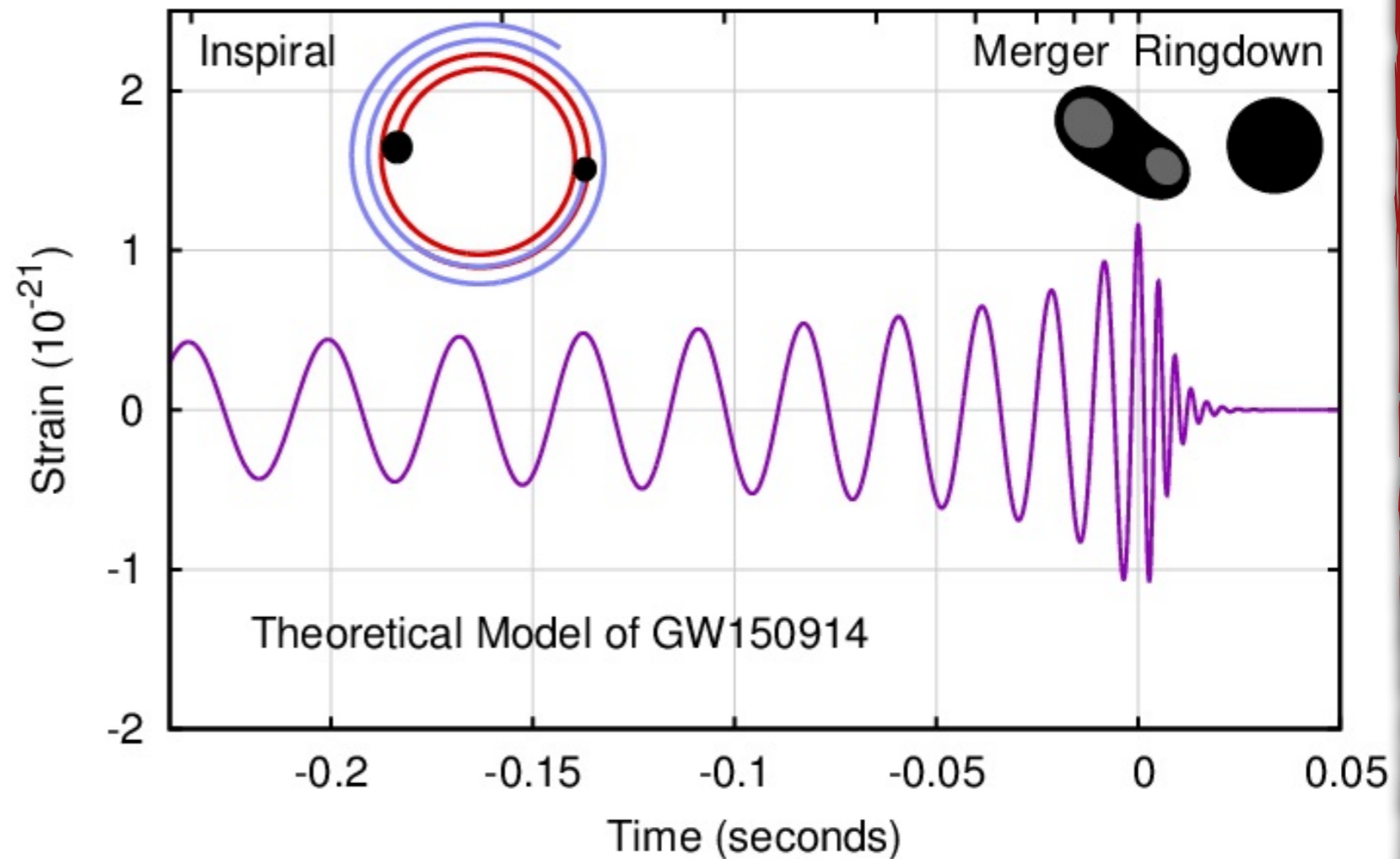
The details of the NS internal structure (and hence its EoS) become important as the orbital separation approaches the size of the bodies



Scattering amplitude

Challenge: two-body problem in General Relativity

Post-newtonian expansion: expansion in v/c



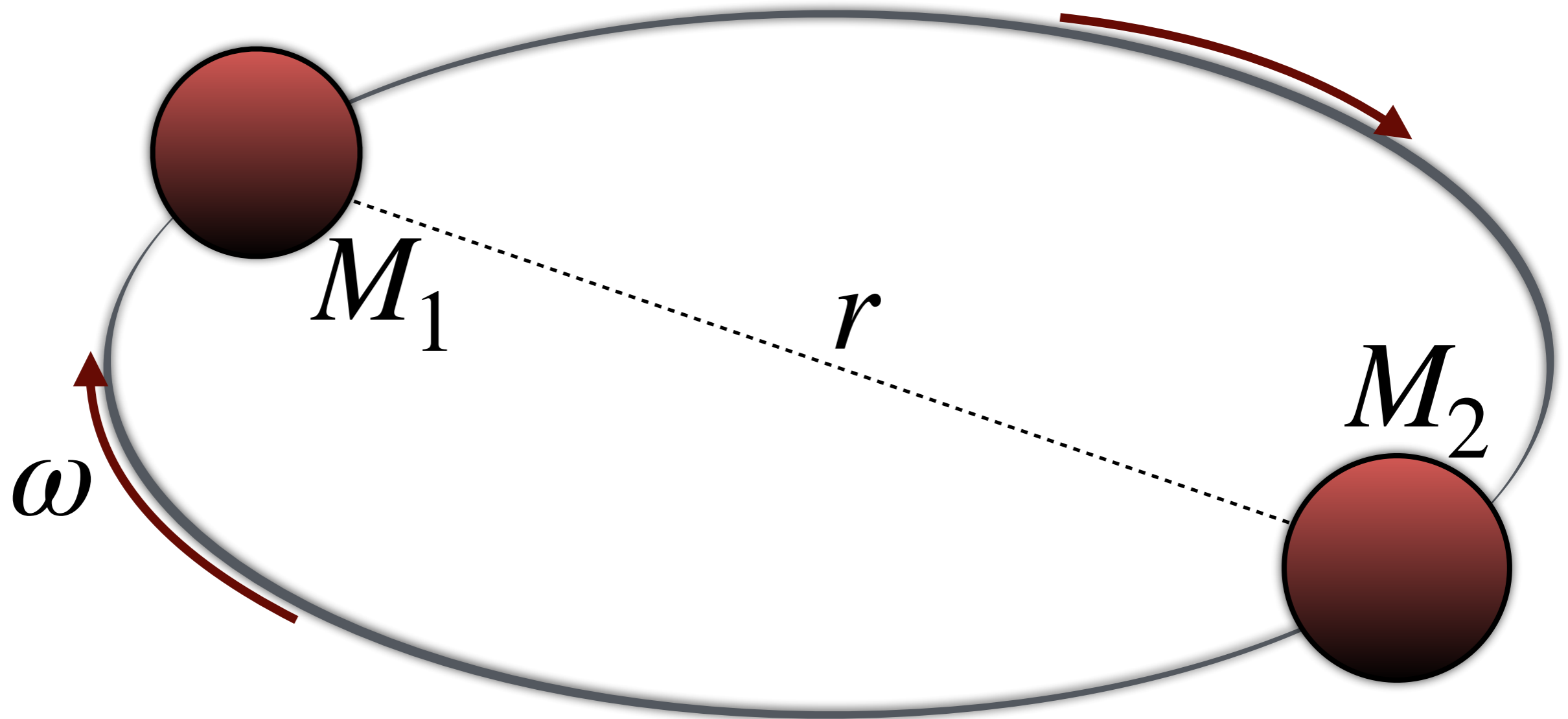
What we can learn
from gravitational waves?



“Gravity is the weakest
force in Nature.”

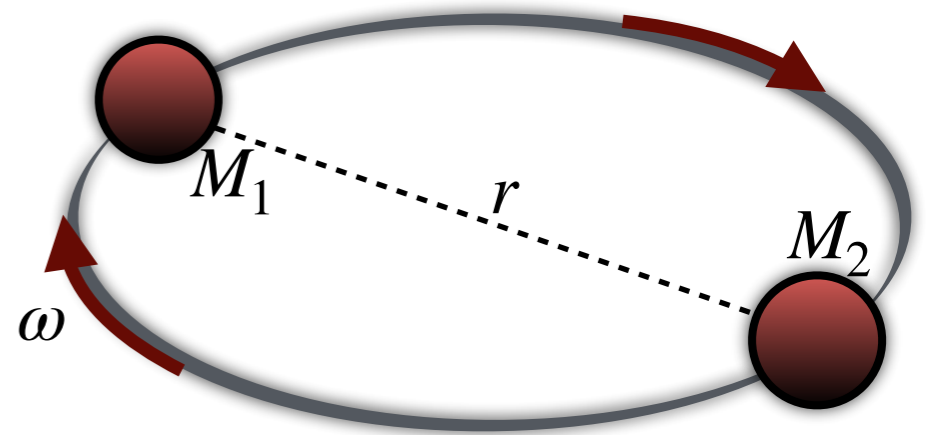
PART II: Phenomenology

PART II: Phenomenology



PART II: Phenomenology

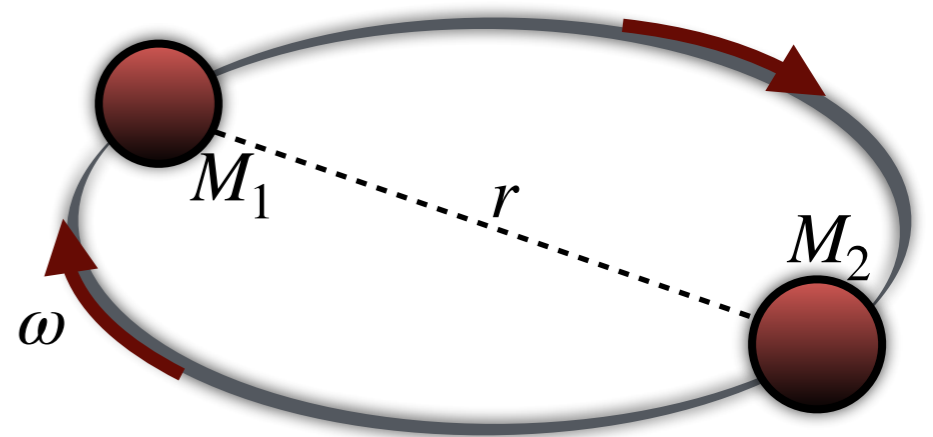
$$\omega = \frac{G_N(M_1 + M_2)}{r^3}$$



PART II: Phenomenology

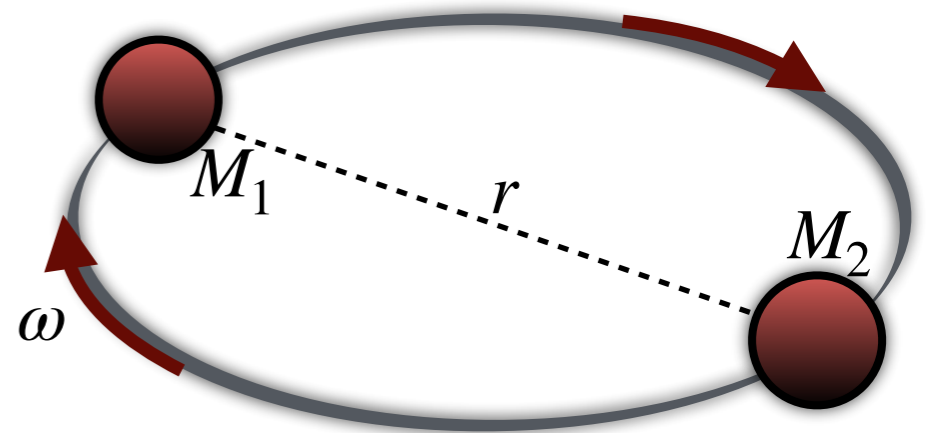
$$\omega = \frac{G_N(M_1 + M_2)}{r^3}$$

$$\frac{dE_{\text{tot}}}{dt} = -\mathcal{P}_{\text{GW}}$$



PART II: Phenomenology

$$\omega = \frac{G_N(M_1 + M_2)}{r^3}$$

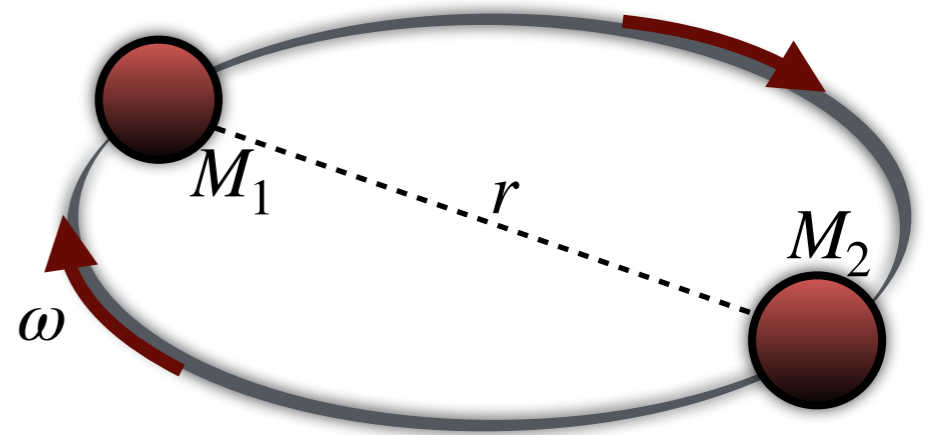


$$\frac{dE_{\text{tot}}}{dt} = -\mathcal{P}_{\text{GW}}$$

$$\mathcal{P}_{\text{GW}} = \frac{32G_N\mu^2\omega^6 r^4}{5}$$

$$E_{\text{tot}} = -\frac{G_N\mu(M_1 + M_2)}{r} + \frac{1}{2}\mu r^2\omega^2$$

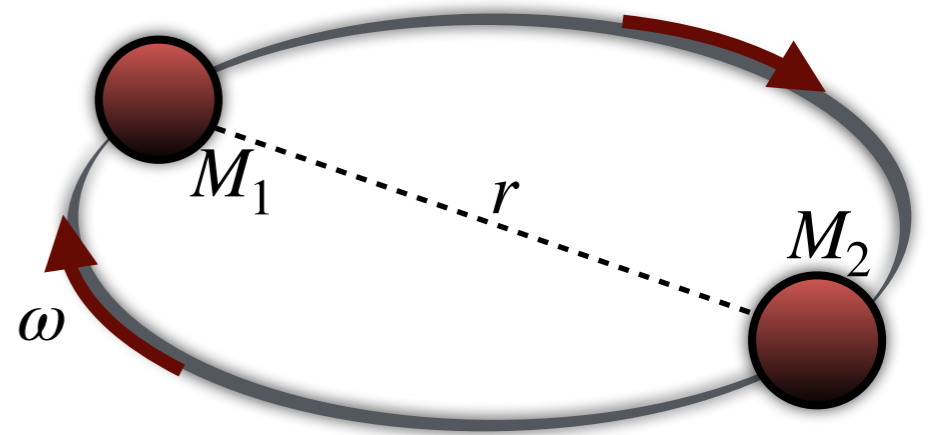
PART II: Phenomenology



$$\frac{d\omega}{dt} = \frac{96}{5} (G_N M_C)^{5/3} \omega^{11/3}$$

PART II: Phenomenology

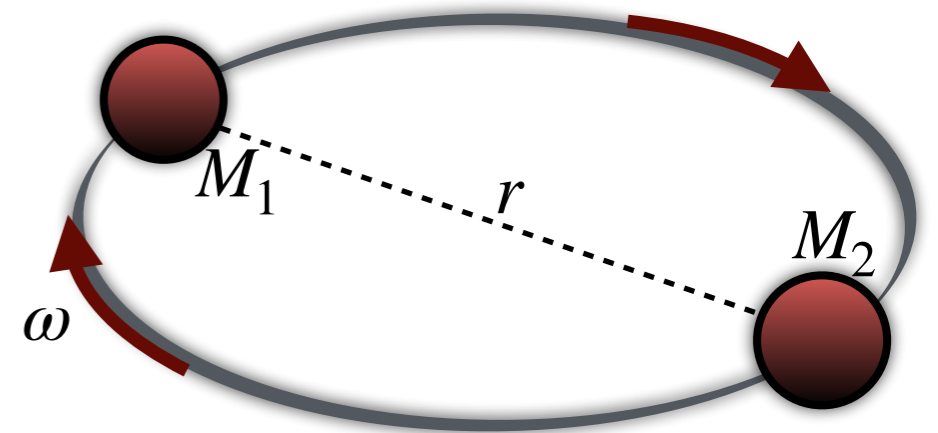
$$M_C \equiv \mu^{3/5} (M_1 + M_2)^{2/5} = \frac{(M_1 M_2)^{3/5}}{(M_1 + M_2)^{1/5}}$$



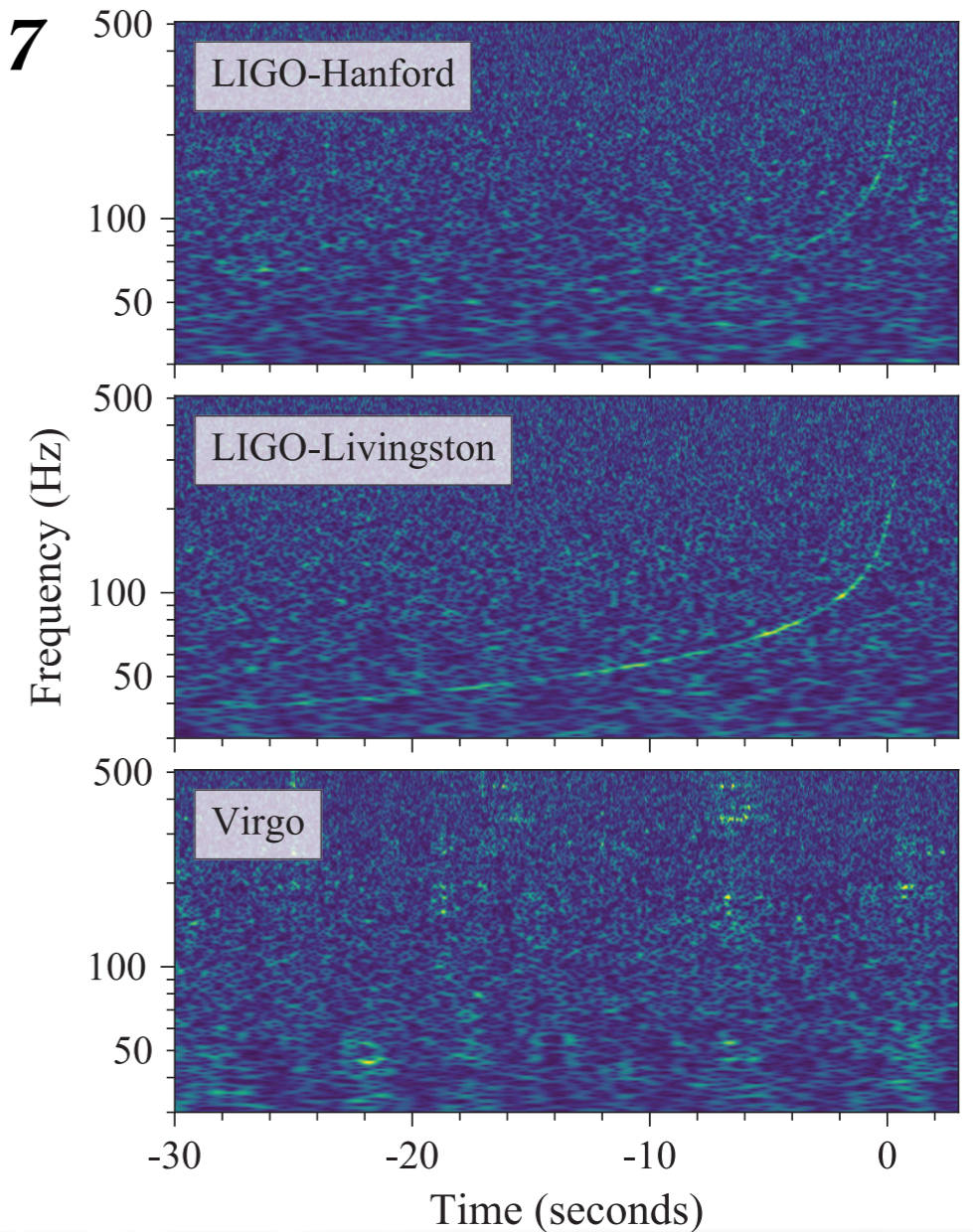
$$\frac{d\omega}{dt} = \frac{96}{5} (G_N M_C)^{5/3} \omega^{11/3}$$

PART II: Phenomenology

$$\frac{d\omega}{dt} = \frac{96}{5} (G_N M_C)^{5/3} \omega^{11/3}$$

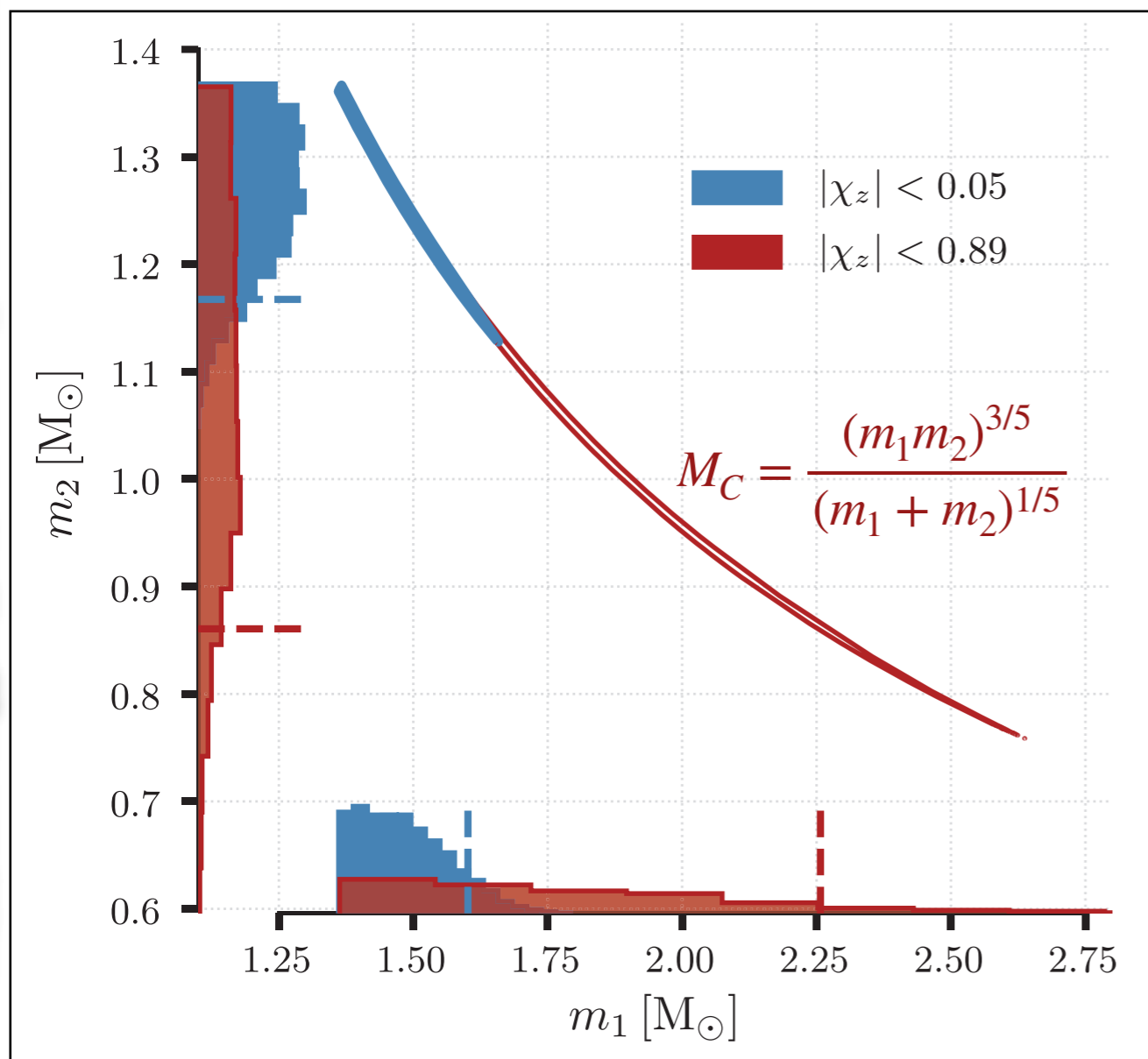
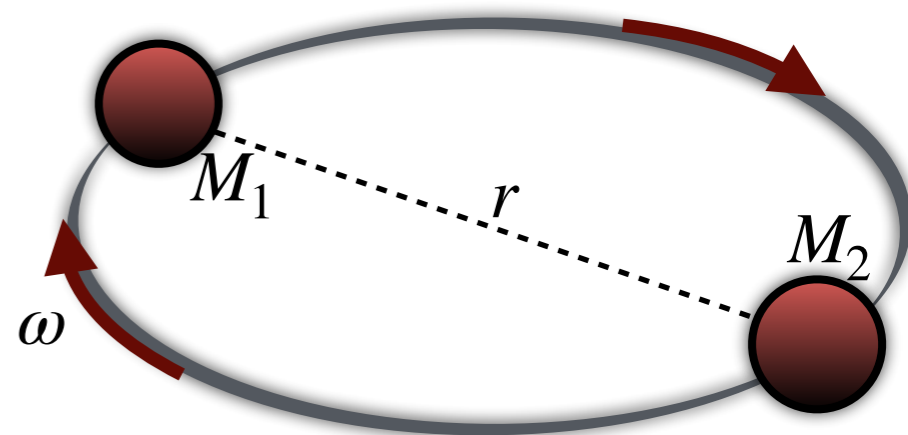


GW170817

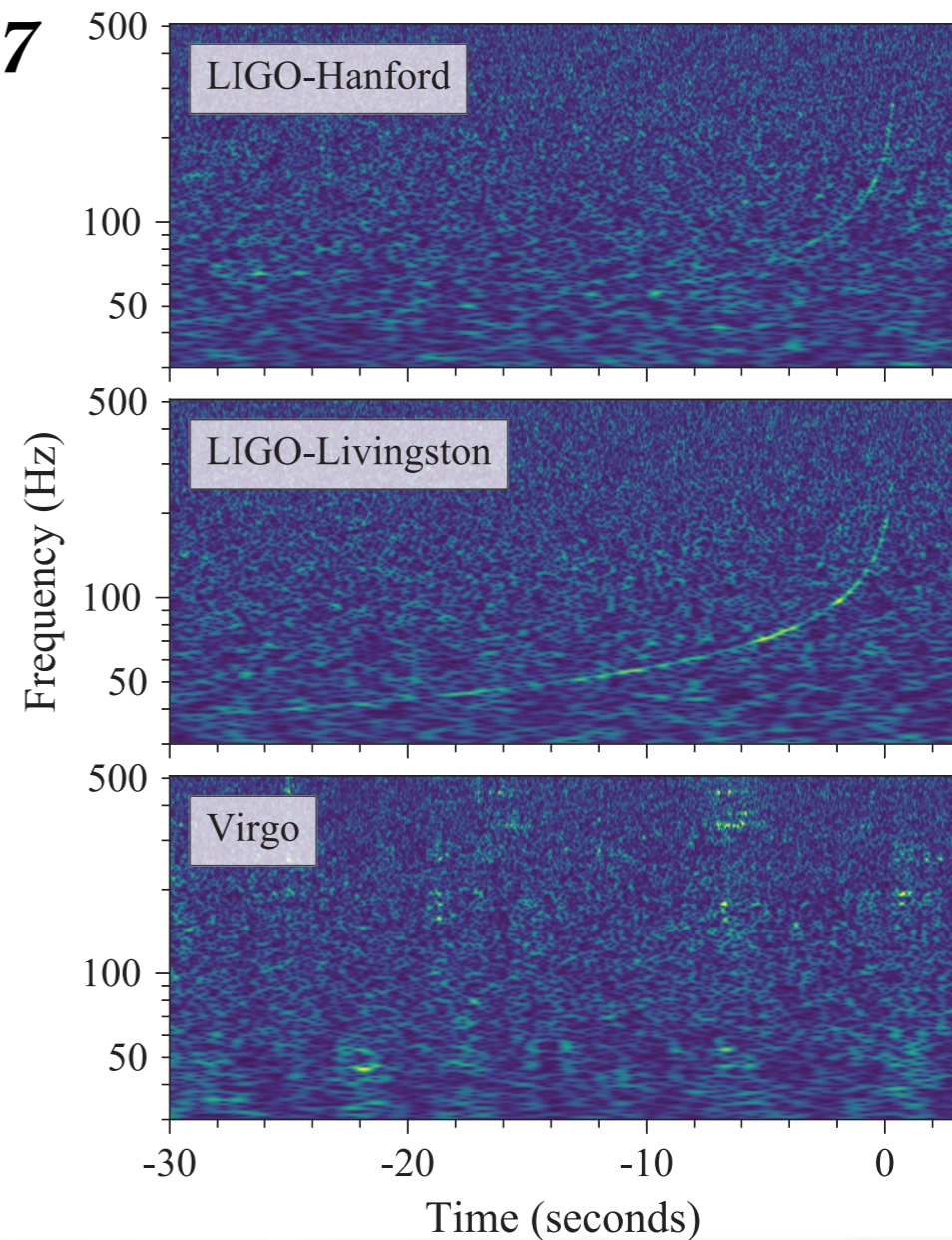


PART II: Phenomenology

$$\frac{d\omega}{dt} = \frac{96}{5} (G_N M_C)^{5/3} \omega^{11/3}$$

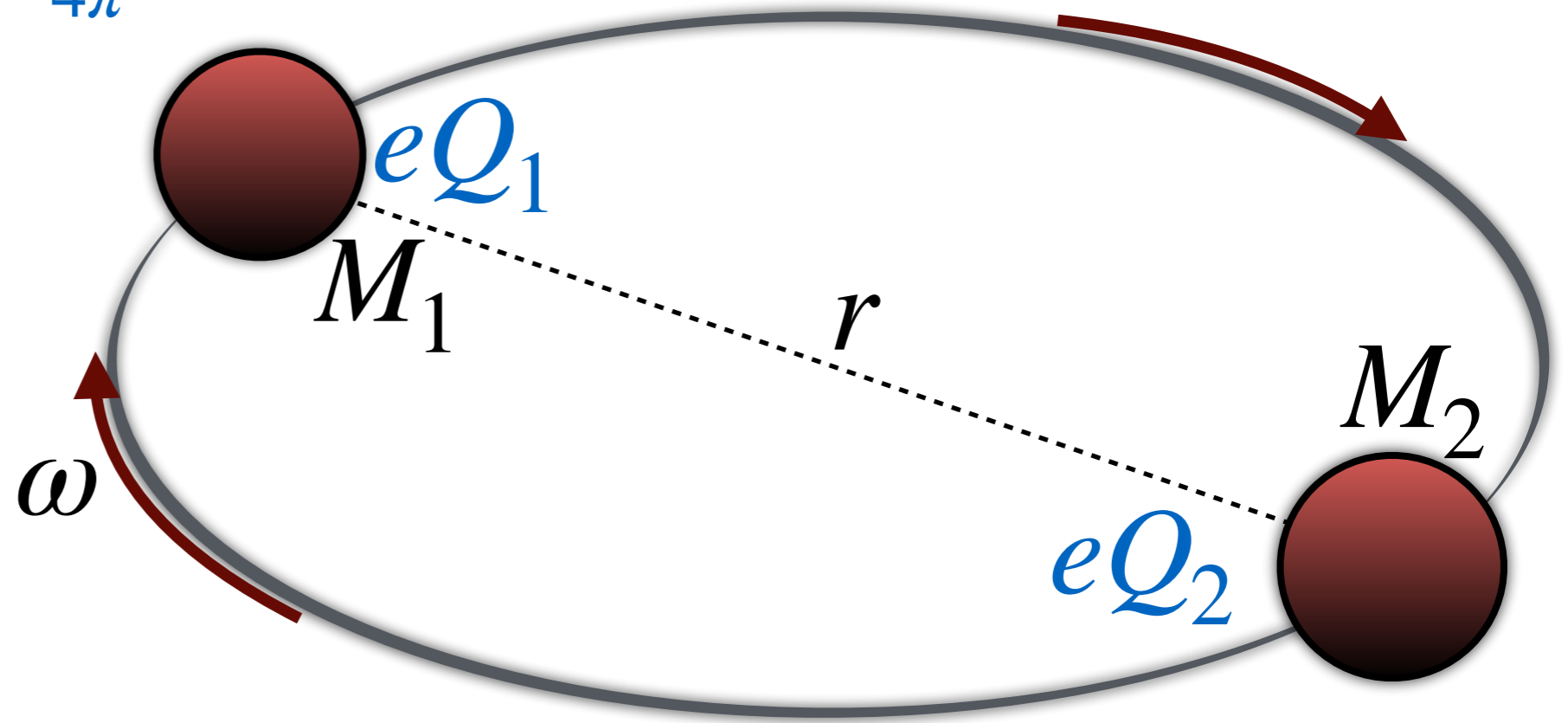


GW170817



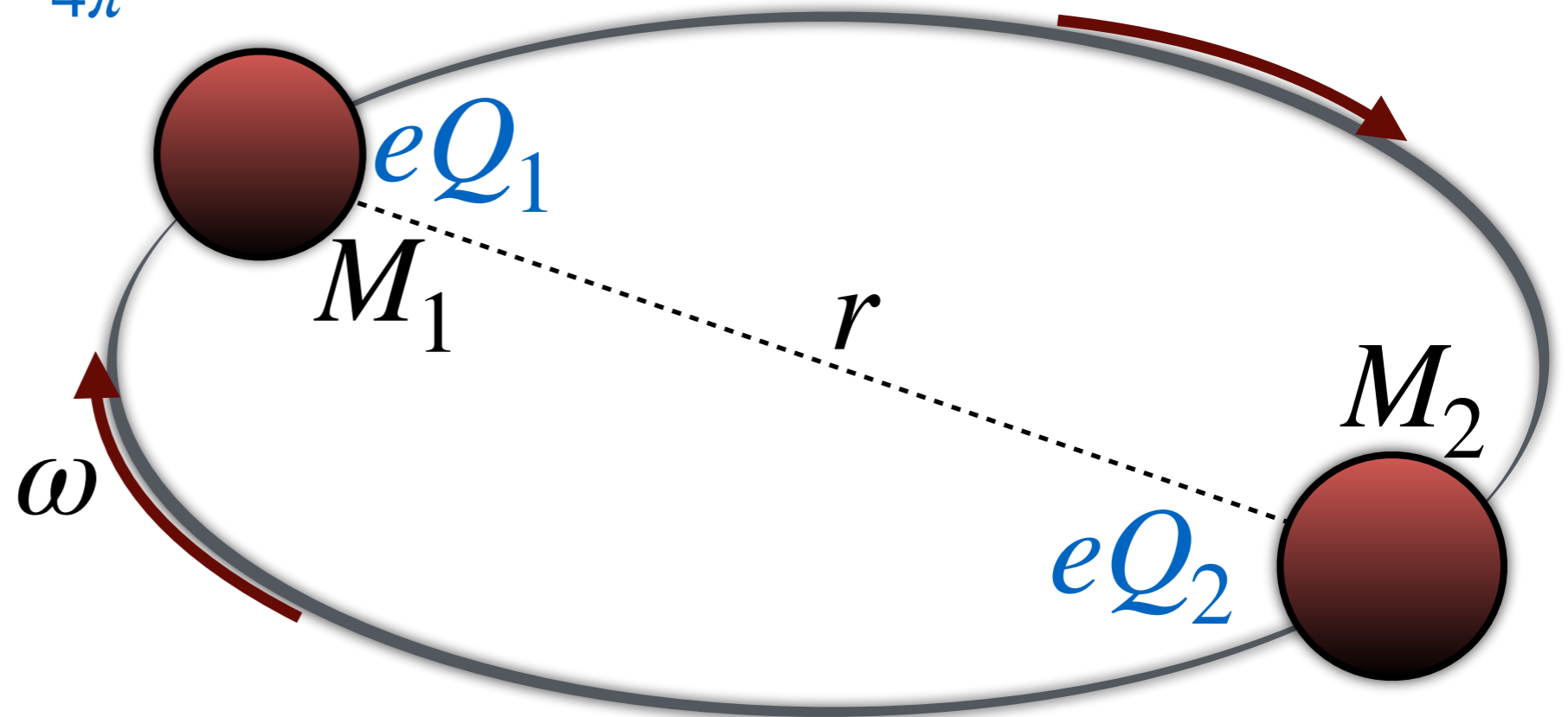
PART II: Phenomenology

$$\alpha \equiv \frac{e^2}{4\pi}$$



PART II: Phenomenology

$$\alpha \equiv \frac{e^2}{4\pi}$$



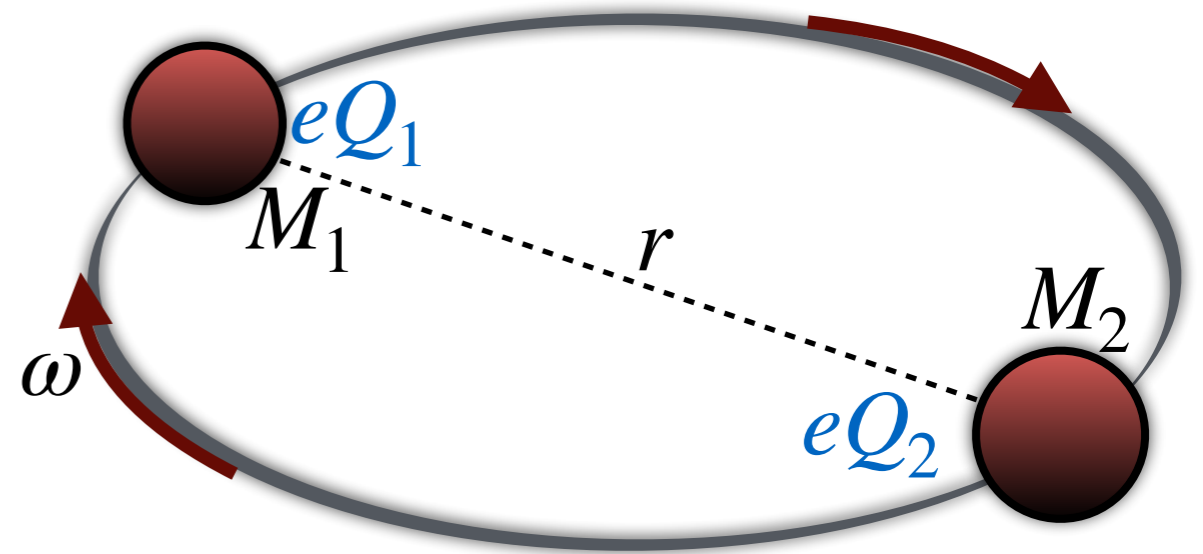
$$V(r) = -\frac{G_N M_1 M_2}{r} + \frac{\alpha Q_1 Q_2}{r}$$

$$\frac{dE_{\text{tot}}}{dt} = -\mathcal{P}_{\text{GW}} - \mathcal{P}_{\text{dark}}$$

$$\mathcal{P}_{\text{dark}} = \frac{2\alpha\gamma^2\omega^4 r^2}{3}$$

PART II: Phenomenology

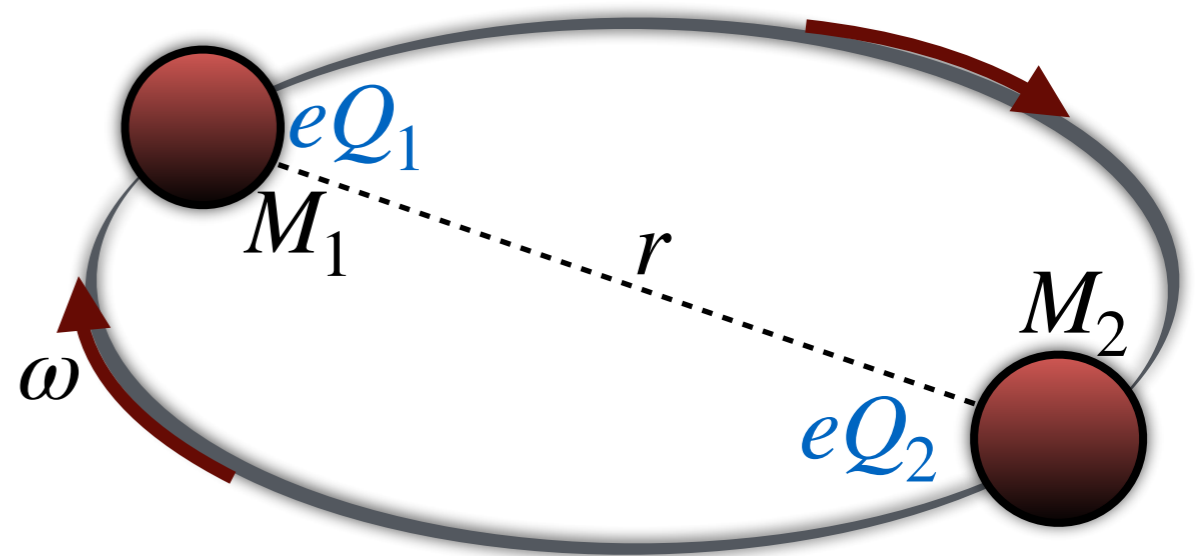
Radiation requires some acceleration of the source



Expansion in multiple of the radiation field at far distance

	Gravitational radiation	Electromagnetic radiation
Monopole	\times Forbidden by conservation of mass	\times Forbidden by conservation of charge
Dipole	\times A dipole depends on the displacement of the center of mass from some fixed point. The velocity of the center of mass is simply given by the total momentum divided by the total mass.	\checkmark (N.B.: unless the charge-to-mass ratio is the same for all the particles)
Quadrupole	\checkmark	\checkmark

PART II: Phenomenology

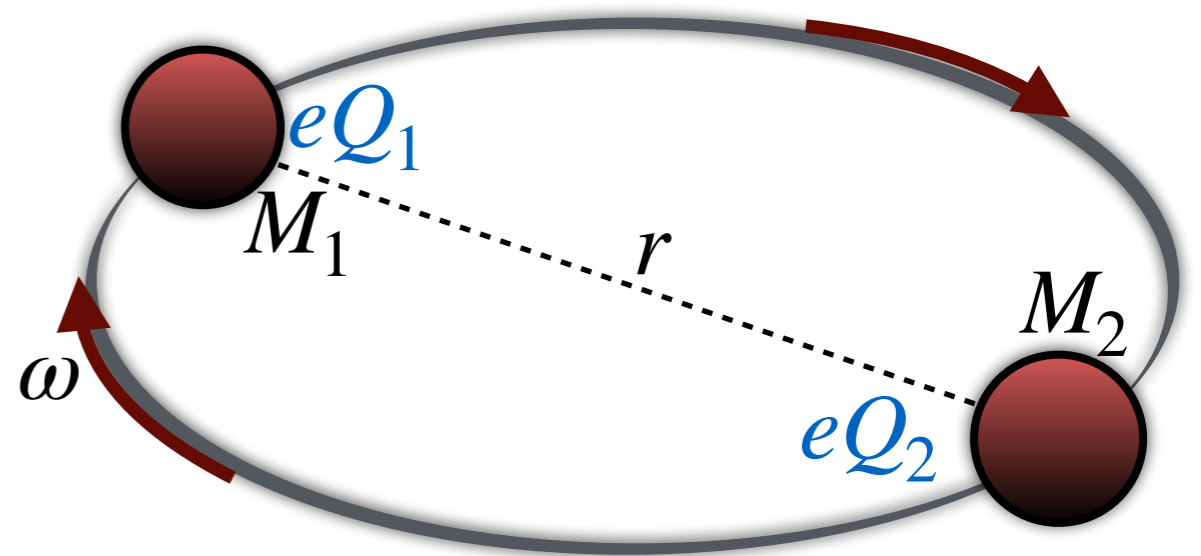


$$\frac{d\omega}{dt} = \frac{96}{5} (G_N M_C)^{5/3} \omega^{11/3} \left(1 - \frac{\alpha Q_1 Q_2}{G_N M_1 M_2} \right)^{2/3} + 2\alpha\mu\omega^3 \left| \frac{Q_1}{M_1} - \frac{Q_2}{M_2} \right|^2$$

PART II: Phenomenology

$$\tilde{\alpha} \equiv \frac{\alpha Q_1 Q_2}{G_N M_1 M_2} = \frac{\alpha Q_1 Q_2}{10^{76}}$$

(with $M_{1,2} = 1.25 M_\odot$)



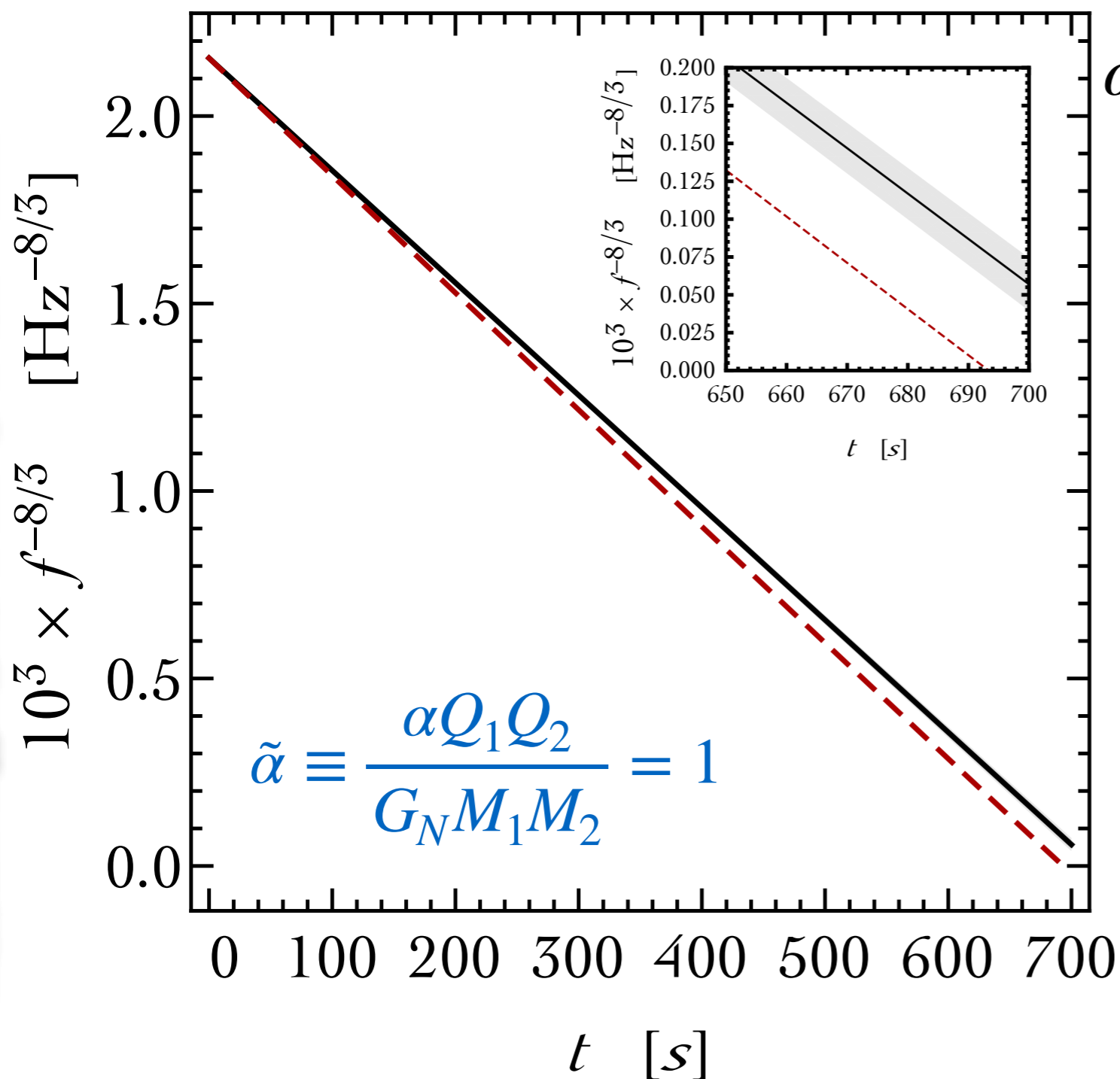
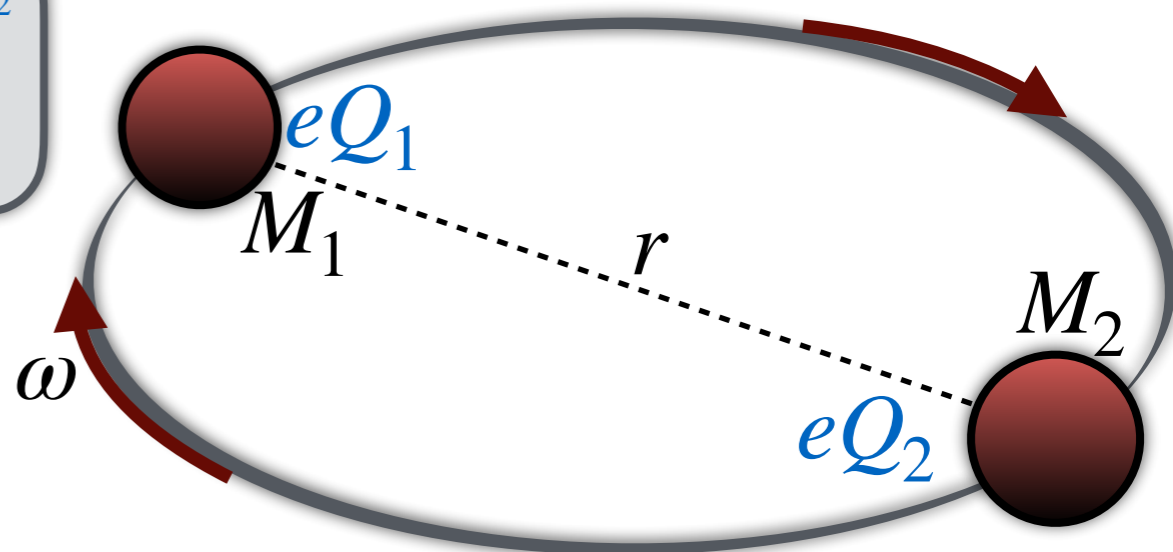
$$\frac{d\omega}{dt} = \frac{96}{5} (G_N M_C)^{5/3} \omega^{11/3} \left(1 - \frac{\alpha Q_1 Q_2}{G_N M_1 M_2} \right)^{2/3} + 2\alpha\mu\omega^3 \left| \frac{Q_1}{M_1} - \frac{Q_2}{M_2} \right|^2$$

**Correction to
quadrupole emission**

Dipole radiation

PART II: Phenomenology

$$\frac{d\omega}{dt} = \frac{96}{5} (G_N M_C)^{5/3} \omega^{11/3} \left(1 - \frac{\alpha Q_1 Q_2}{G_N M_1 M_2} \right)^{2/3} + 2\alpha\mu\omega^3 \left| \frac{Q_1}{M_1} - \frac{Q_2}{M_2} \right|^2$$

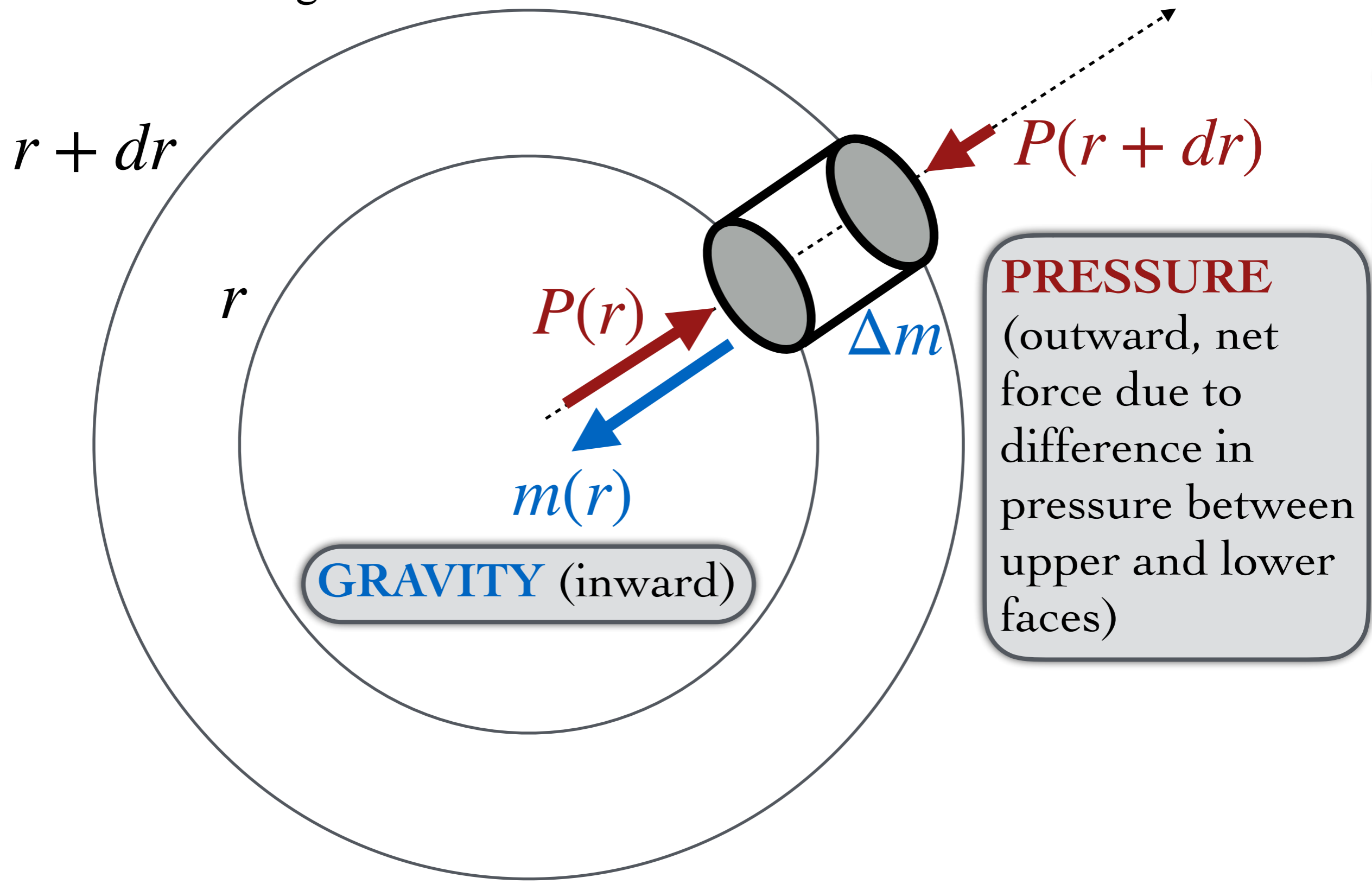


assuming 0.1 % error on M_C

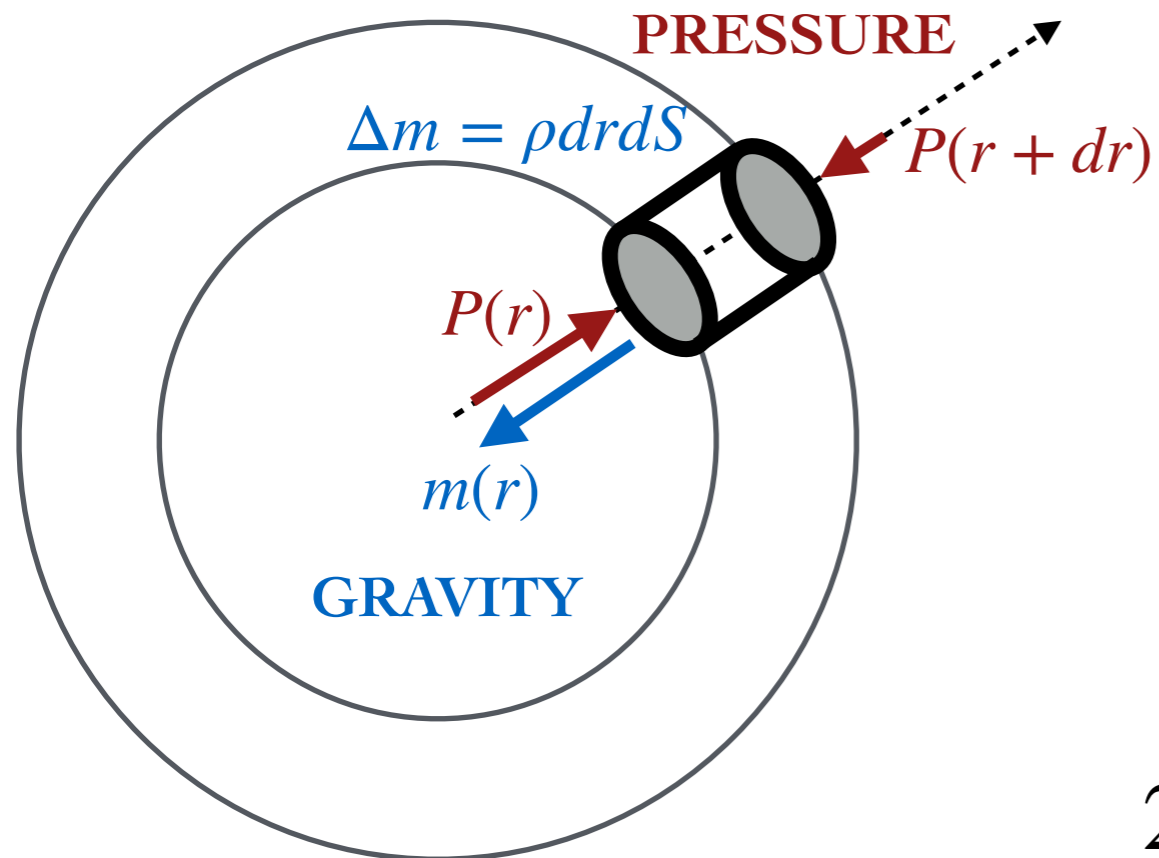
$$\tilde{\alpha} \gtrsim 10^{-2}$$

PART II: Phenomenology

Radial forces acting on the infinitesimal element



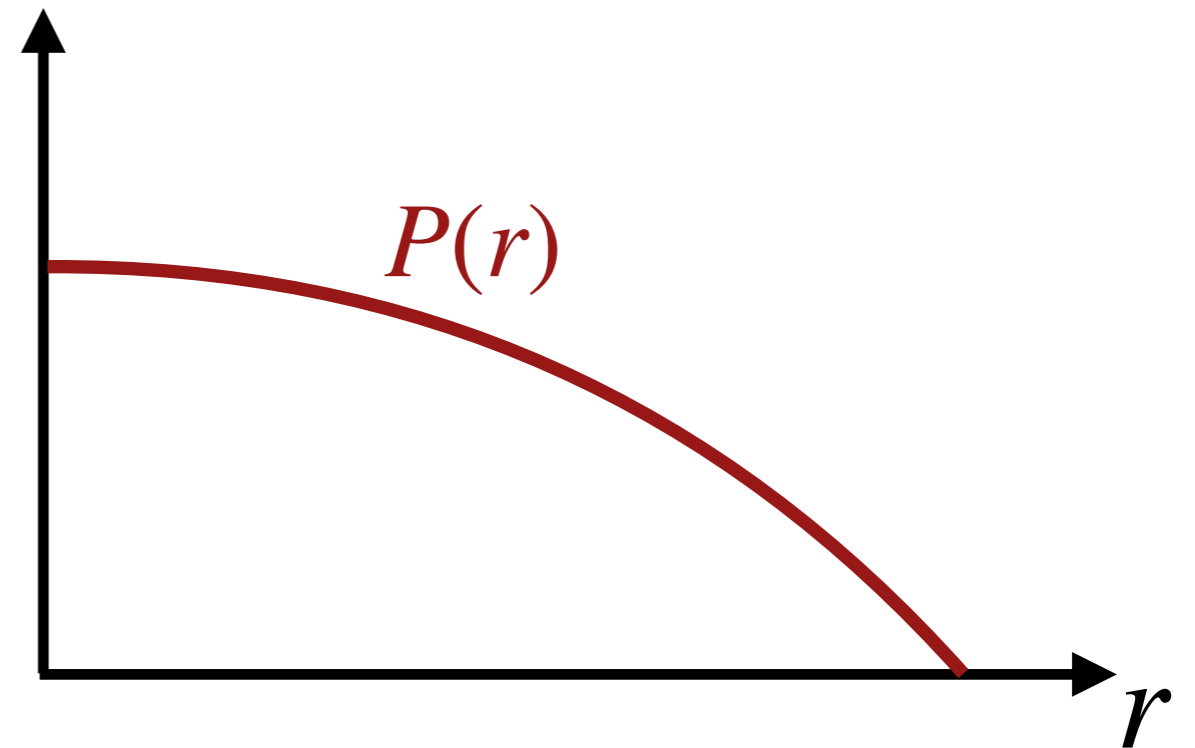
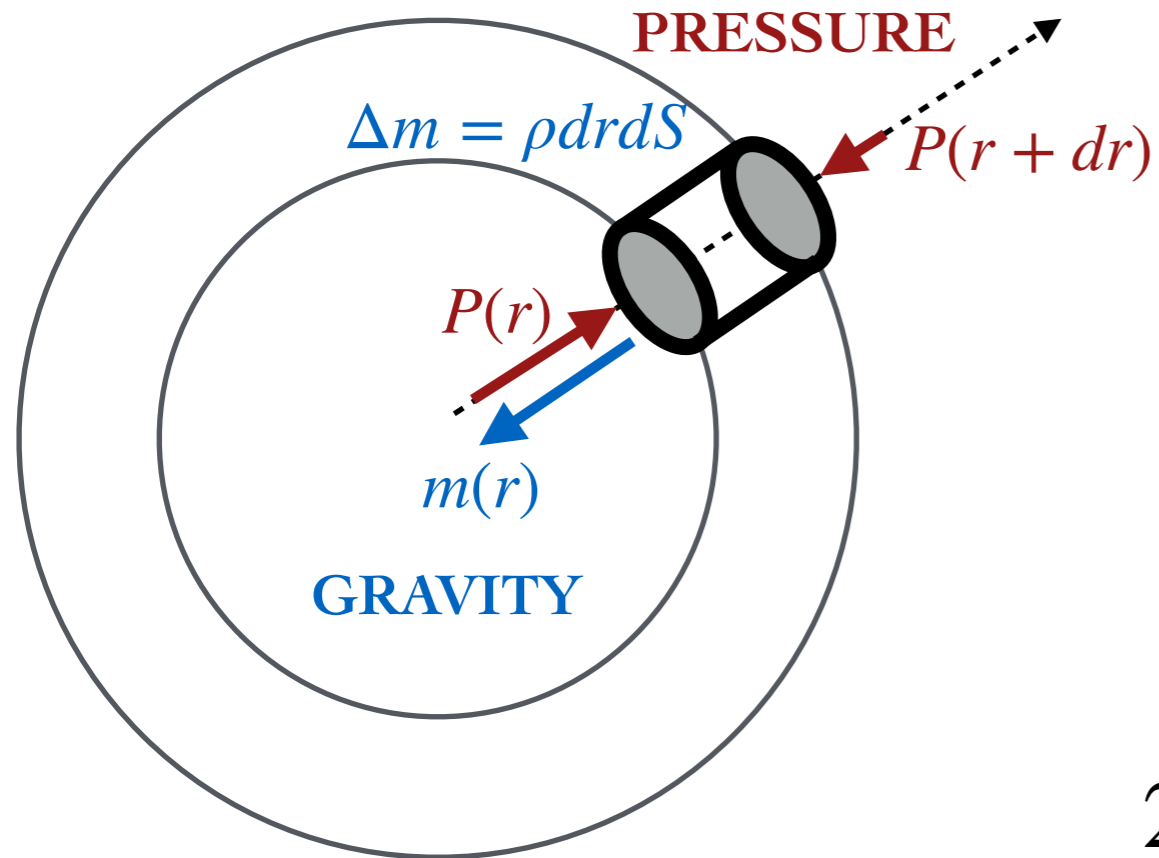
PART II: Phenomenology



2nd equation of stellar structure
(hydrostatic equilibrium)

$$\frac{dP}{dr} = - \frac{G_N \rho m}{r^2}$$

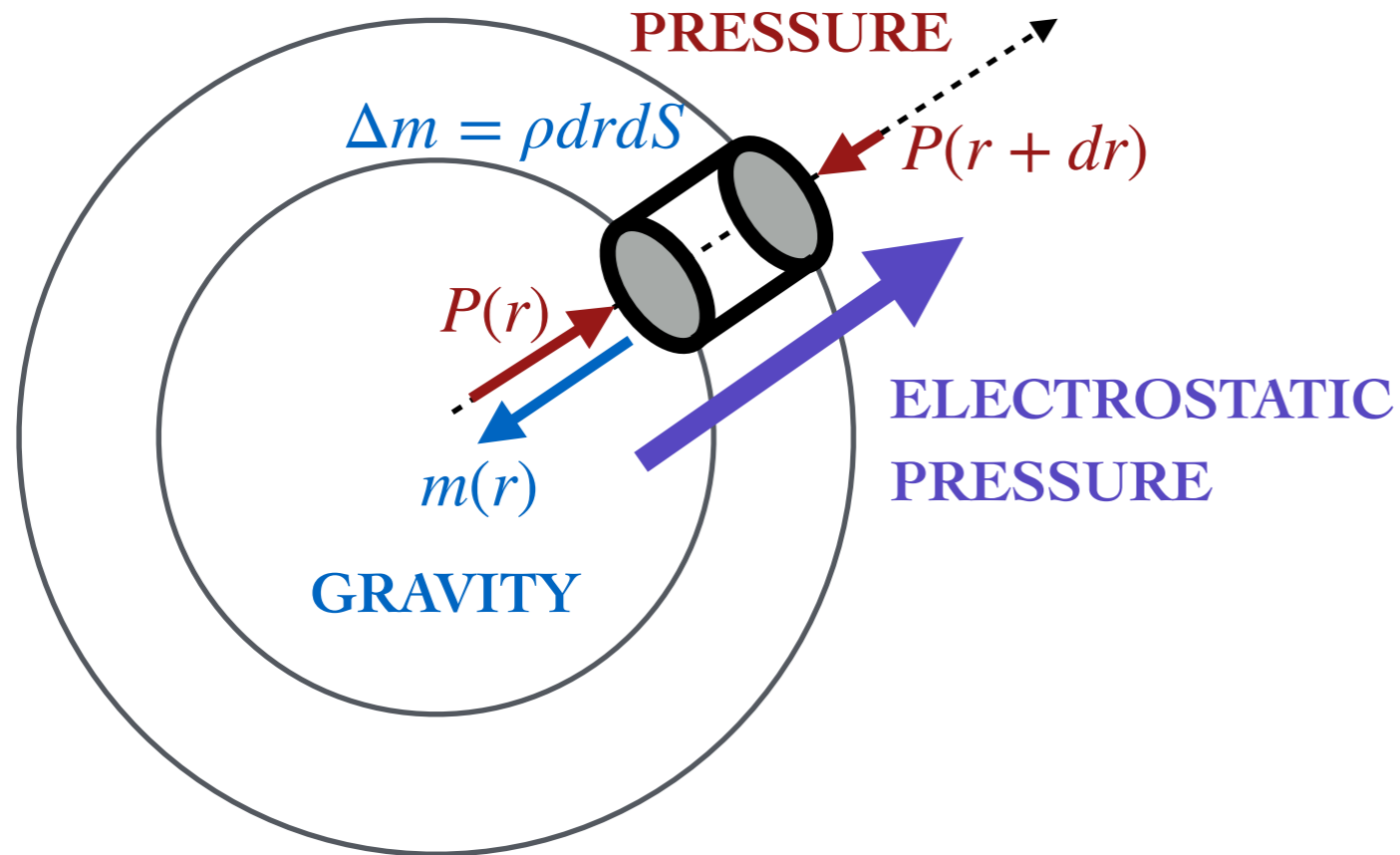
PART II: Phenomenology



2nd stellar structure equation
(hydrostatic equilibrium)

$$\frac{dP}{dr} = - \frac{G_N \rho m}{r^2}$$

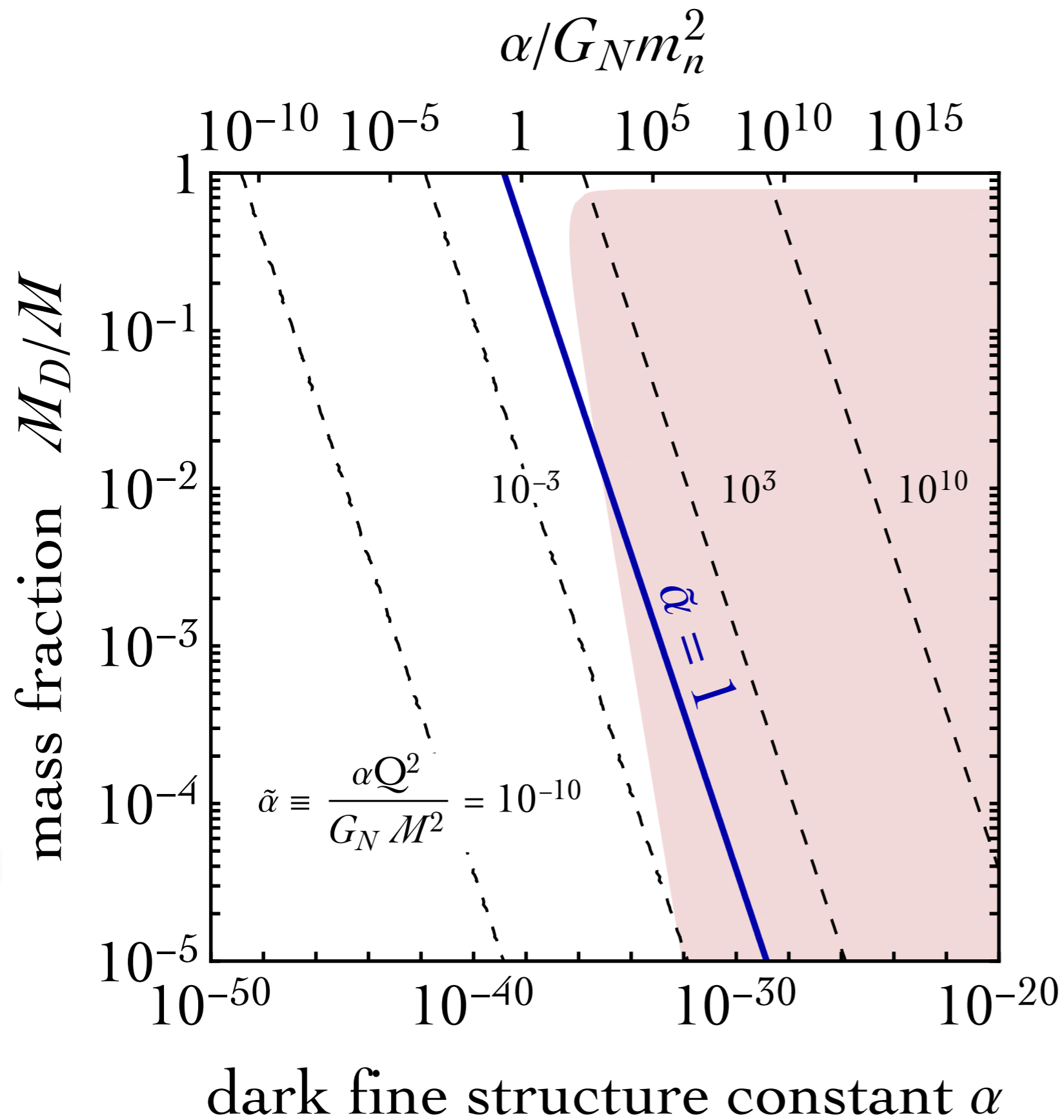
PART II: Phenomenology



$$E(r) = \frac{Q(r)}{4\pi r^2}$$

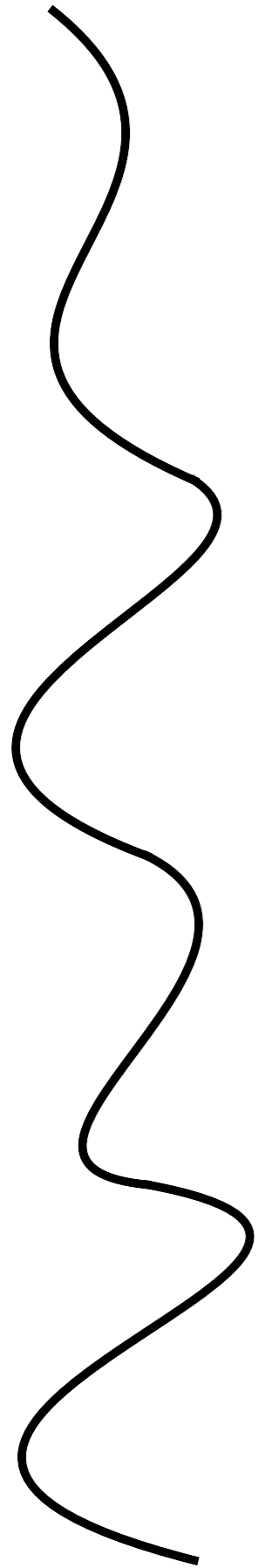
$$\frac{dP}{dr} = - \frac{G_N \rho}{r^2} \left(m - 2\pi r^3 E^2 \right)$$

PART II: Phenomenology



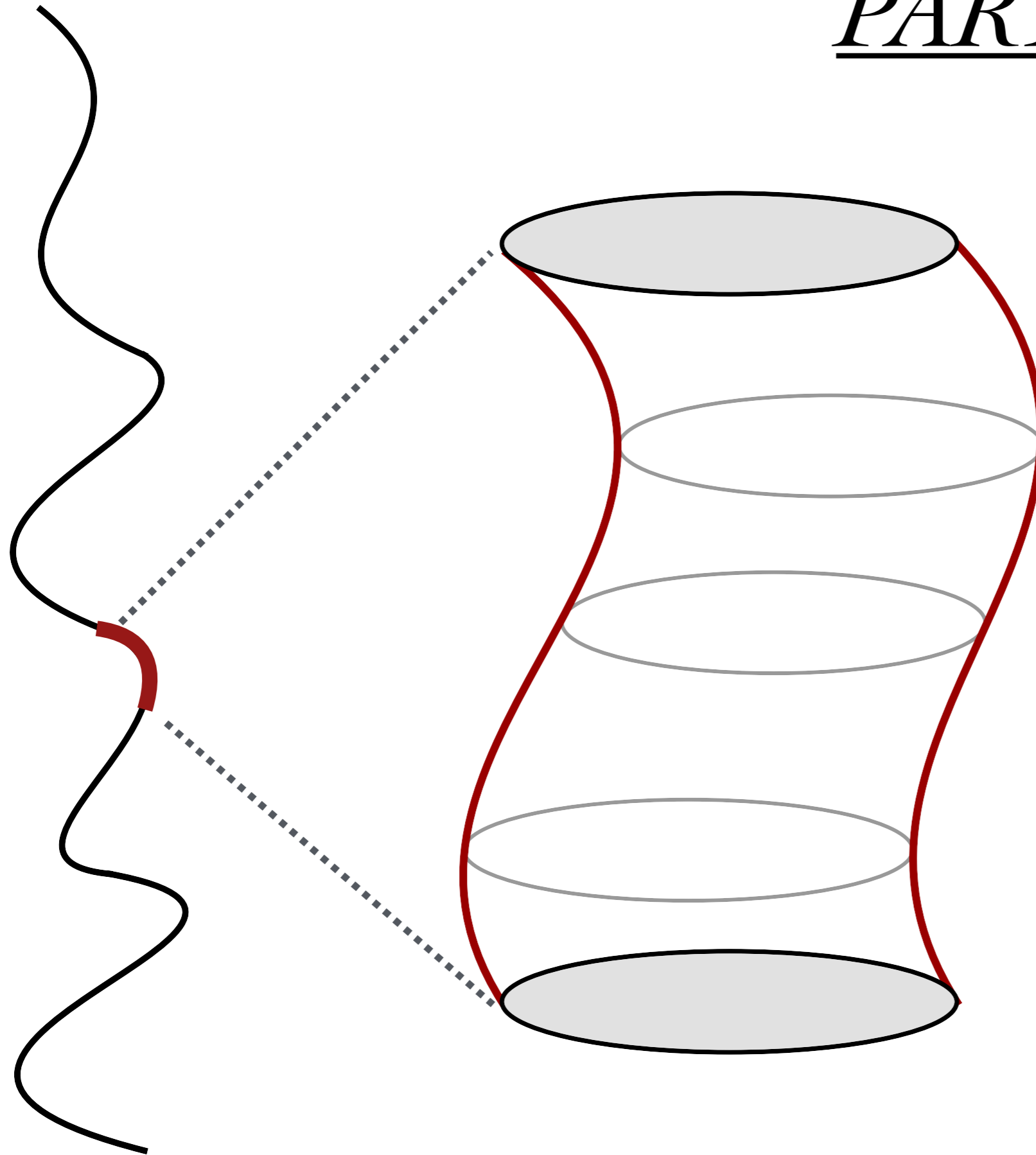
PART III: Theory

PART III: Theory



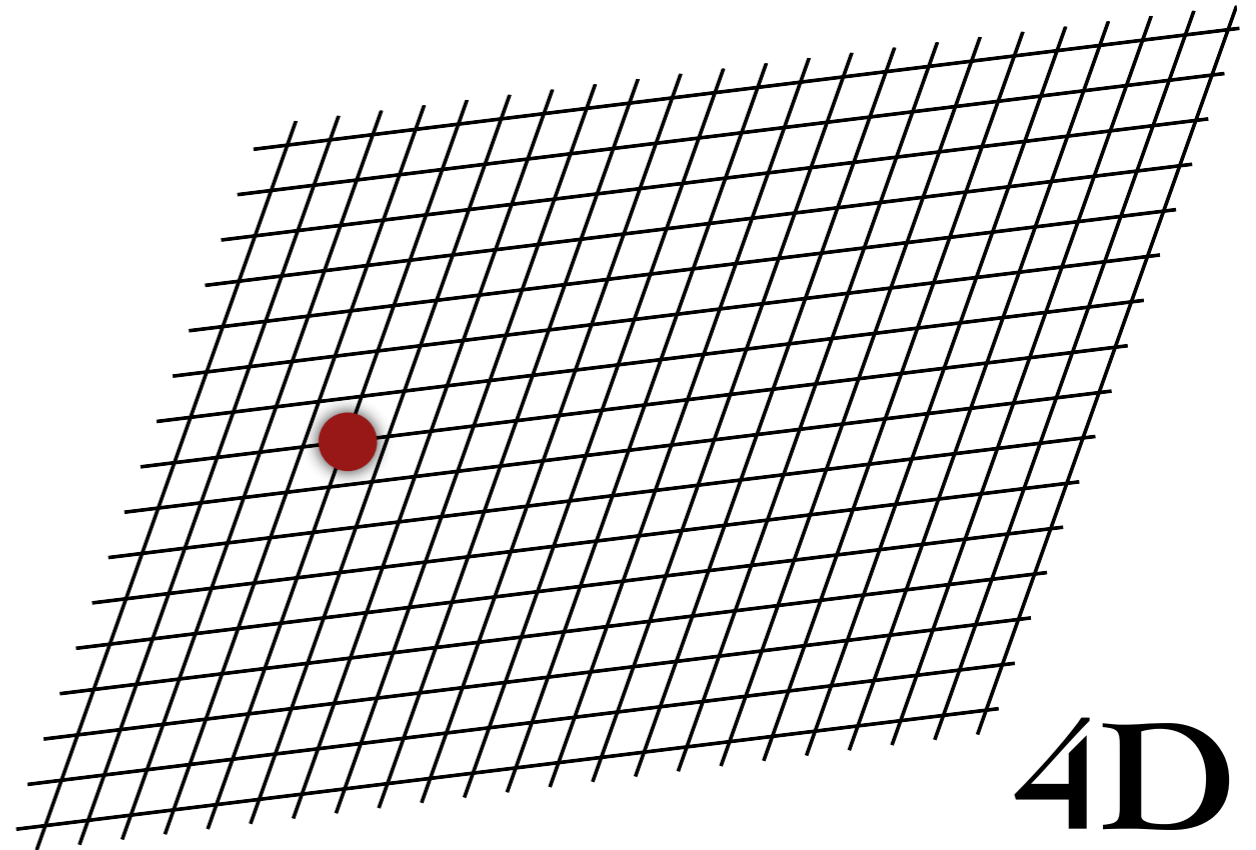
Consider a space that consists of a long, thin tube. Viewed from far distance, the tube looks like a one-dimensional line.

PART III: Theory



Under high magnification, the cylindrical shape becomes apparent. Each point on the line is revealed to be a one-dimensional circle of the tube.

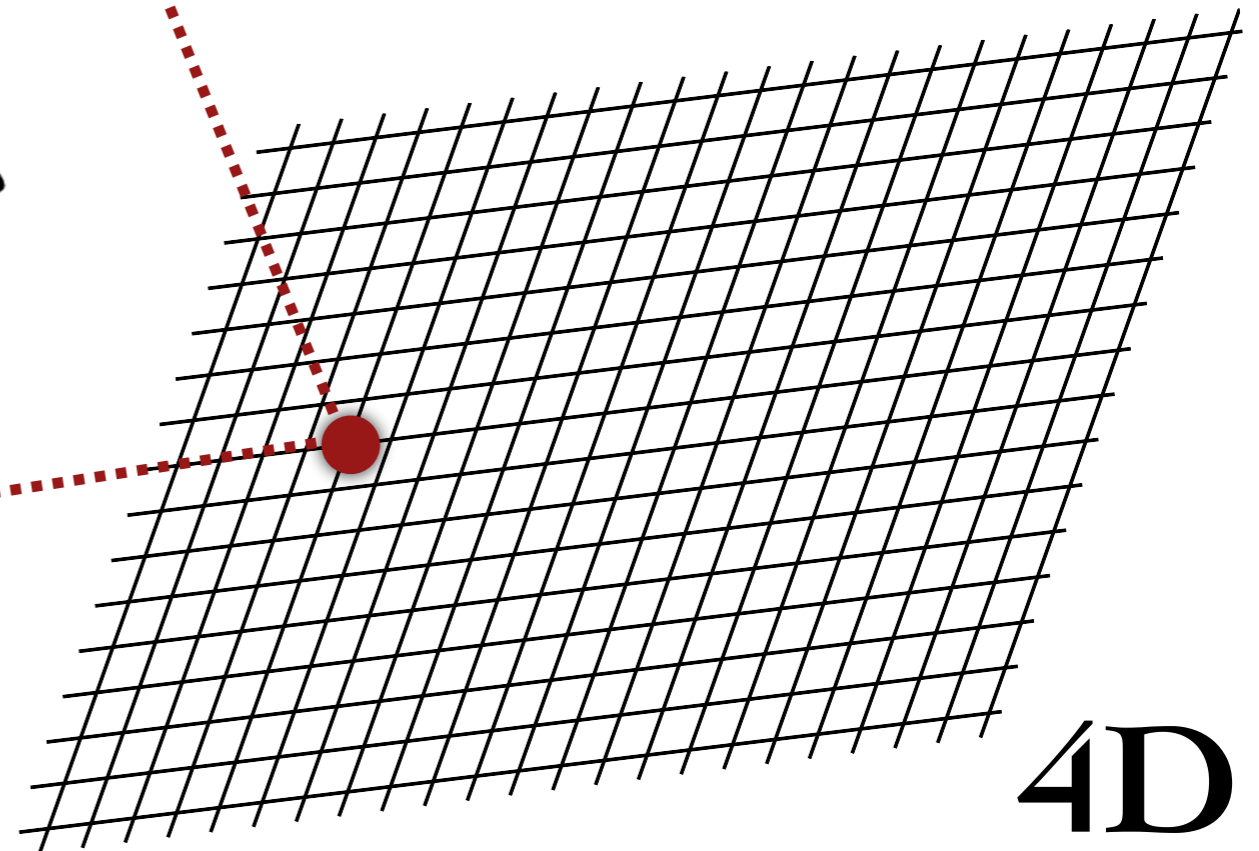
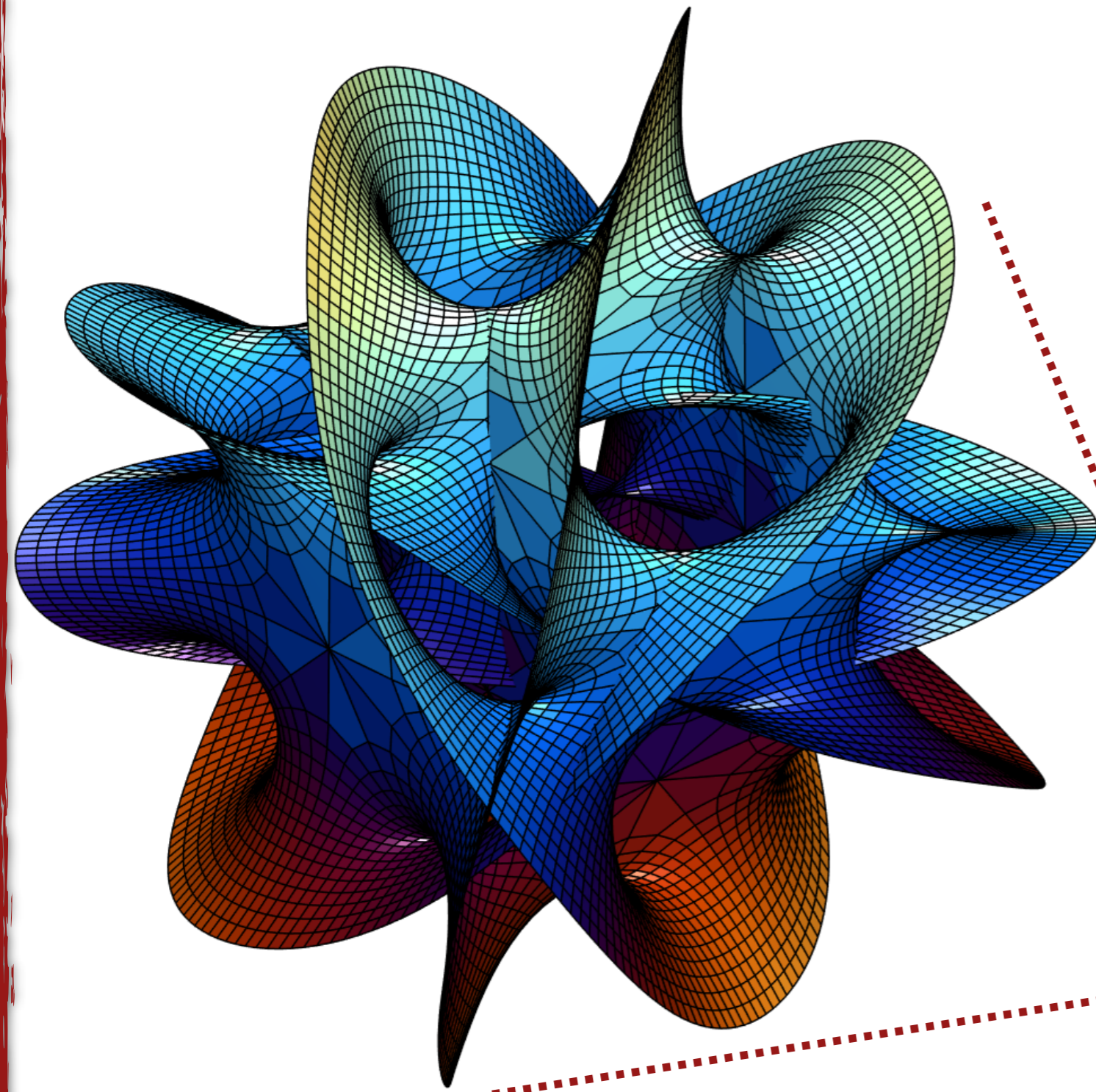
PART III: Theory



4D

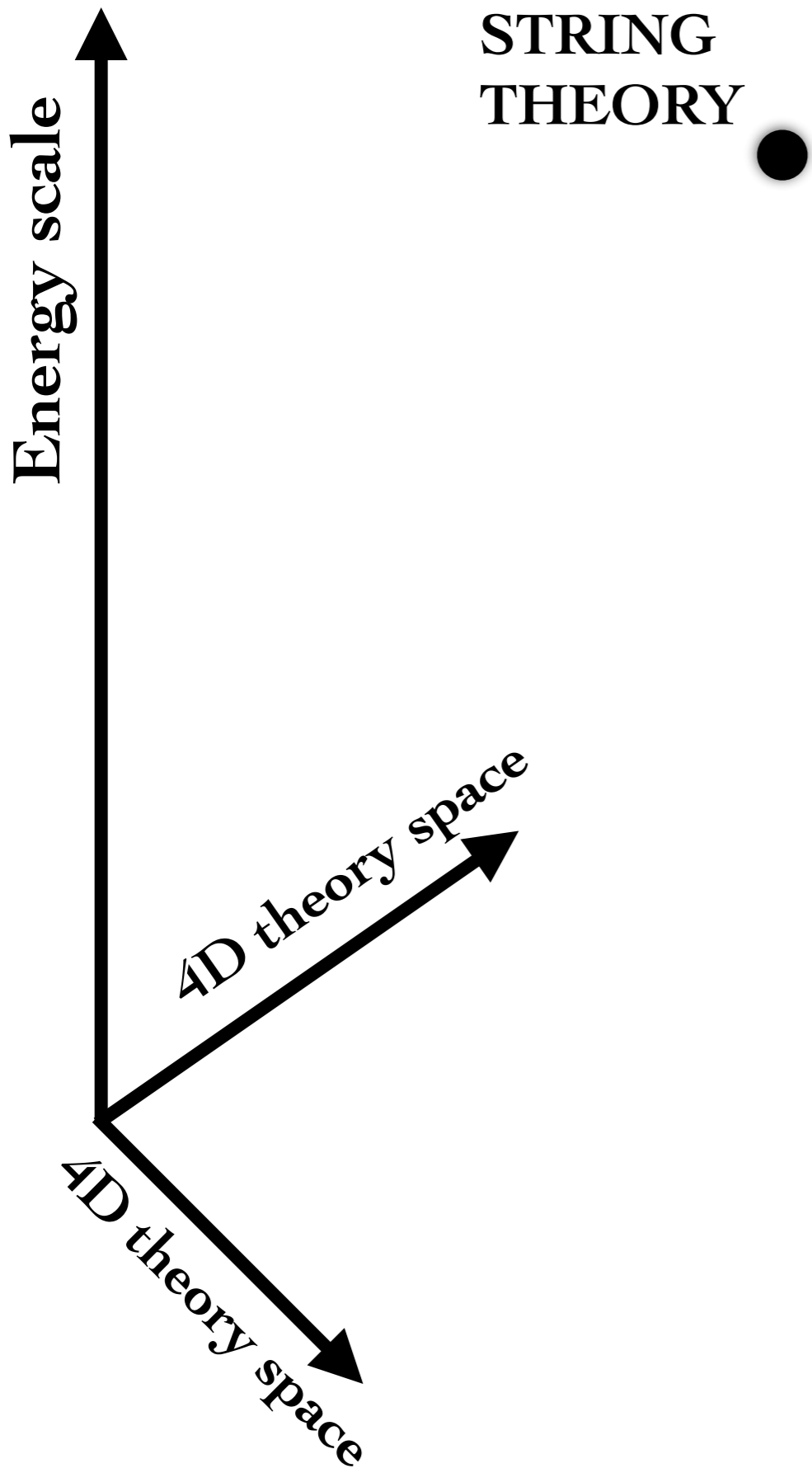
PART III: Theory

Our world has six extra dimensions, every point of our familiar space hides an associated tiny six-dimensional space. The physics that is observed depends on the size and the structure of the manifold.



4D

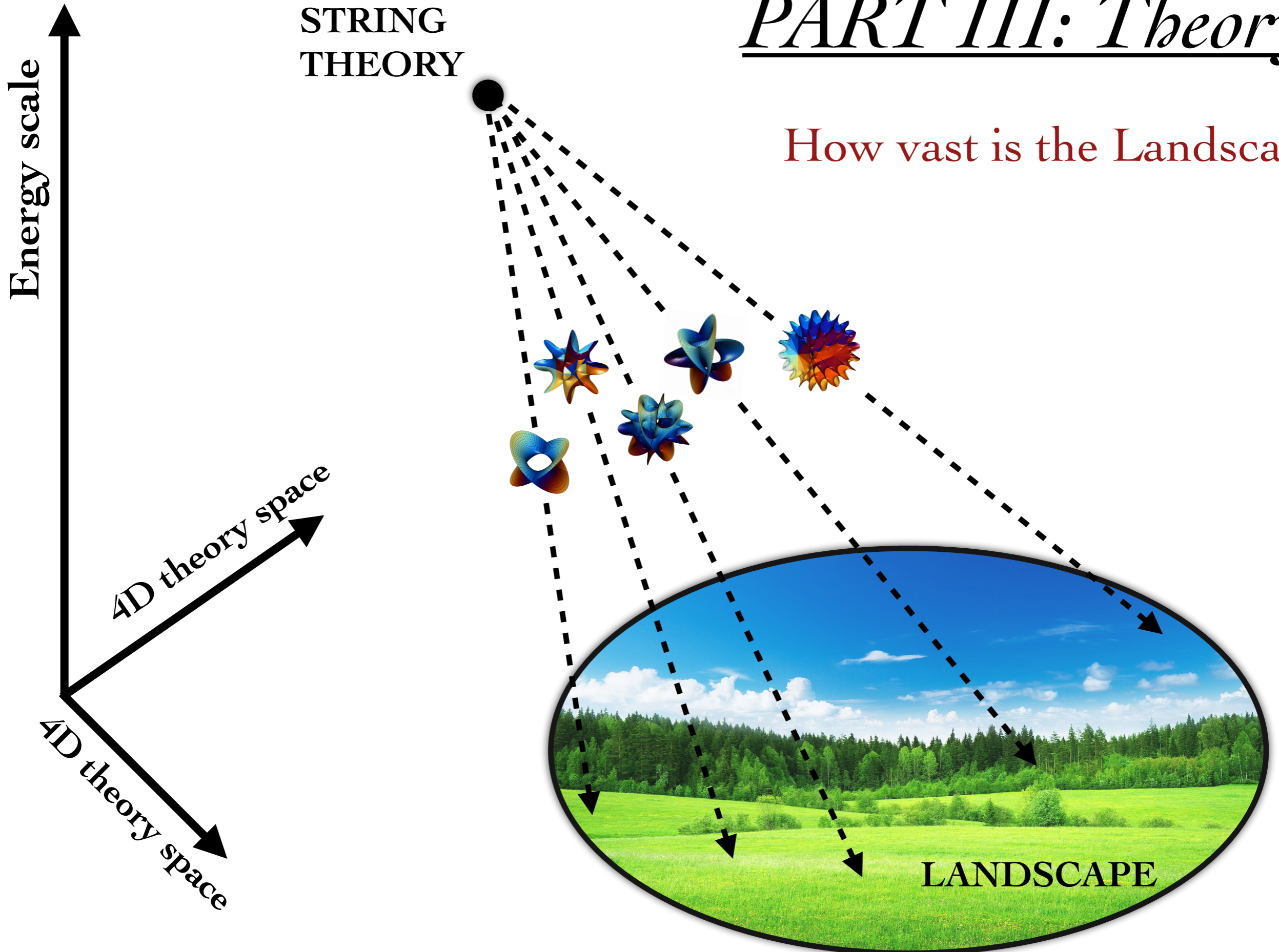
PART III: Theory



PART III: Theory

STRING
THEORY

How vast is the Landscape?



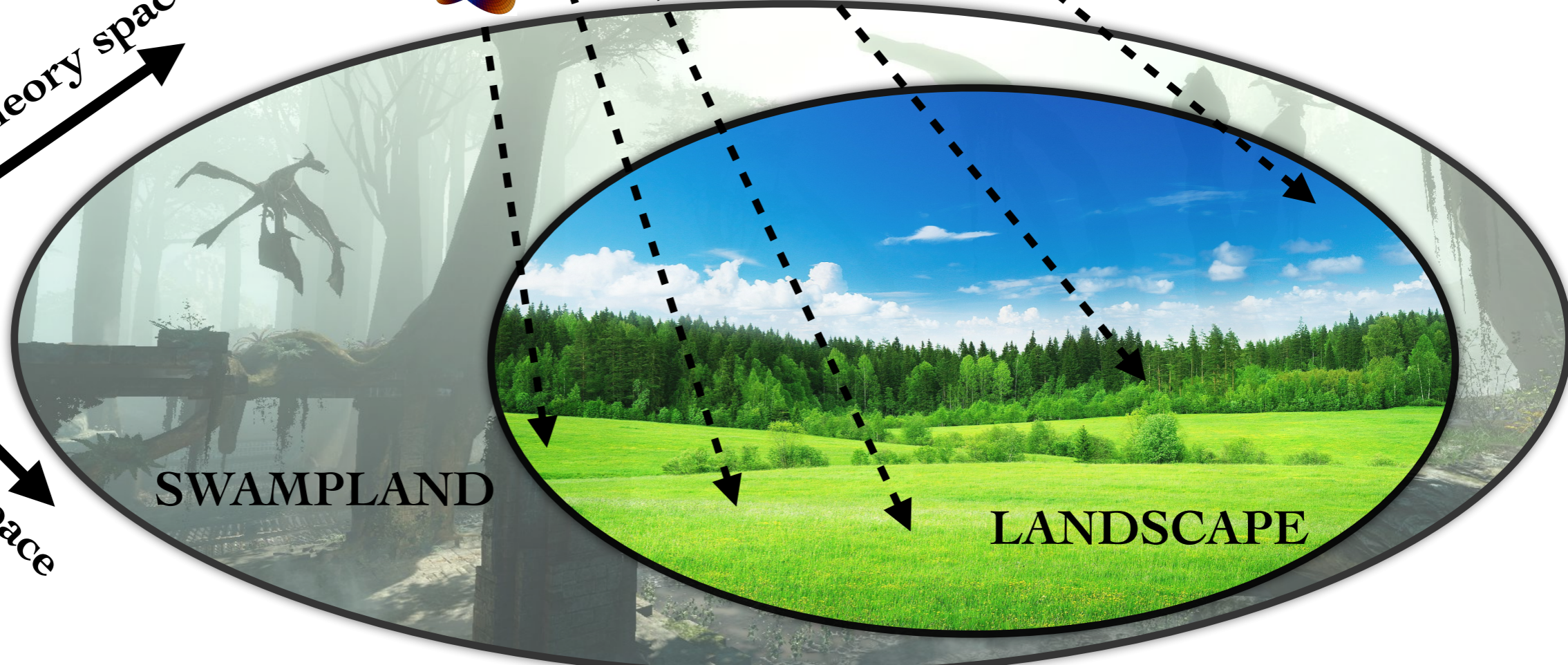
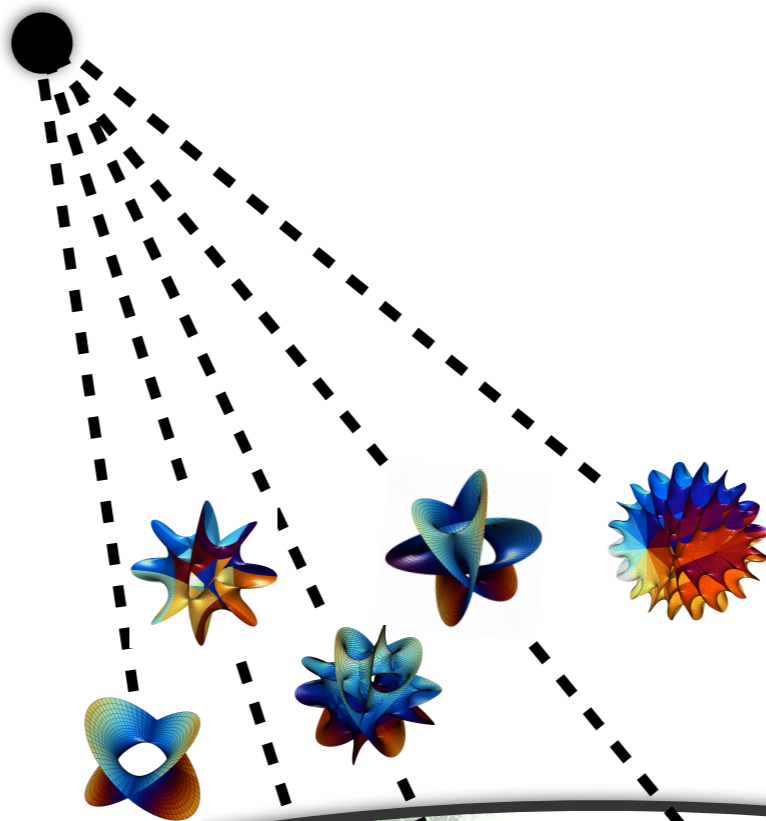
PART III: Theory

STRING
THEORY

Energy scale

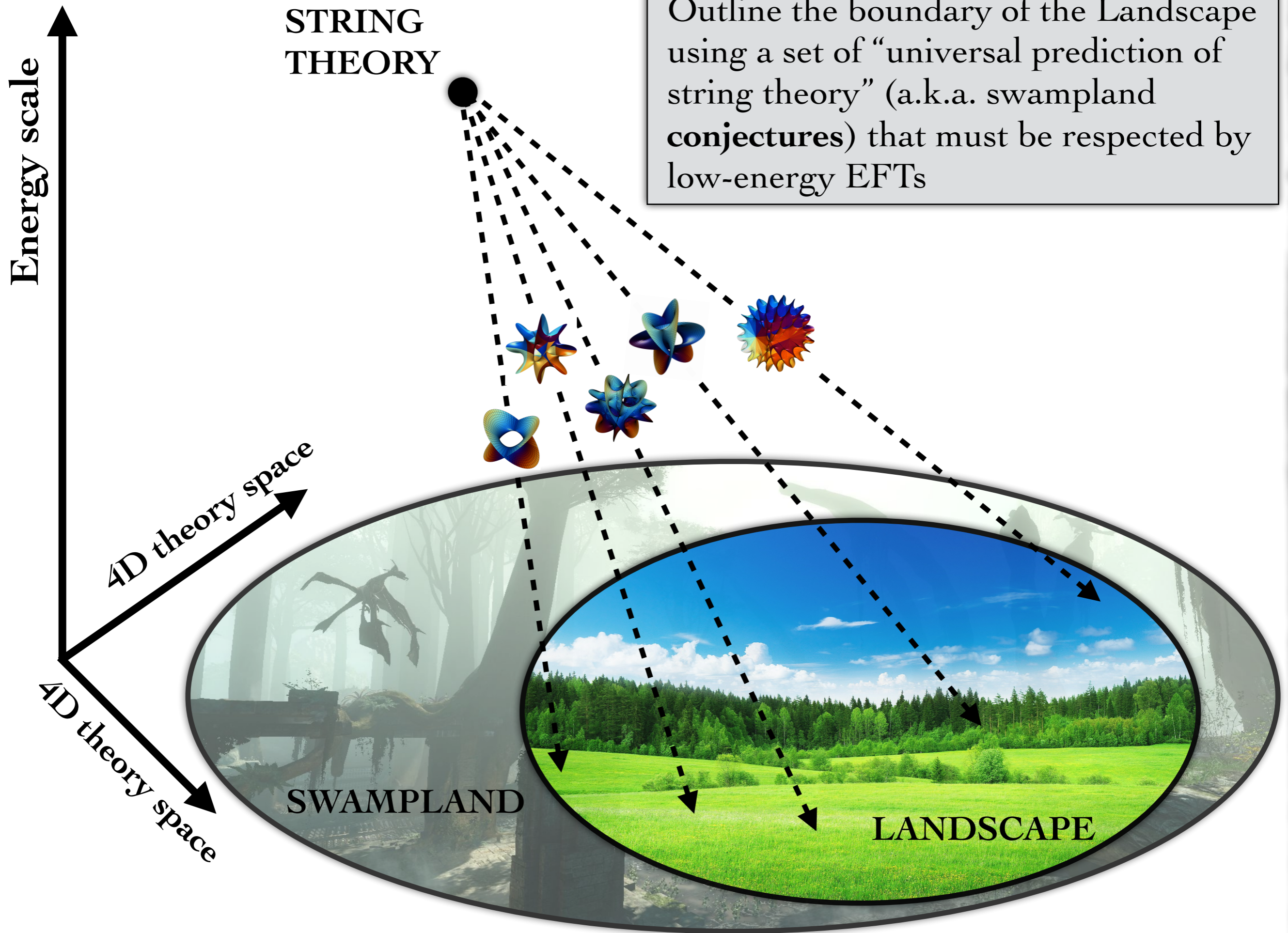
4D theory space

4D theory space

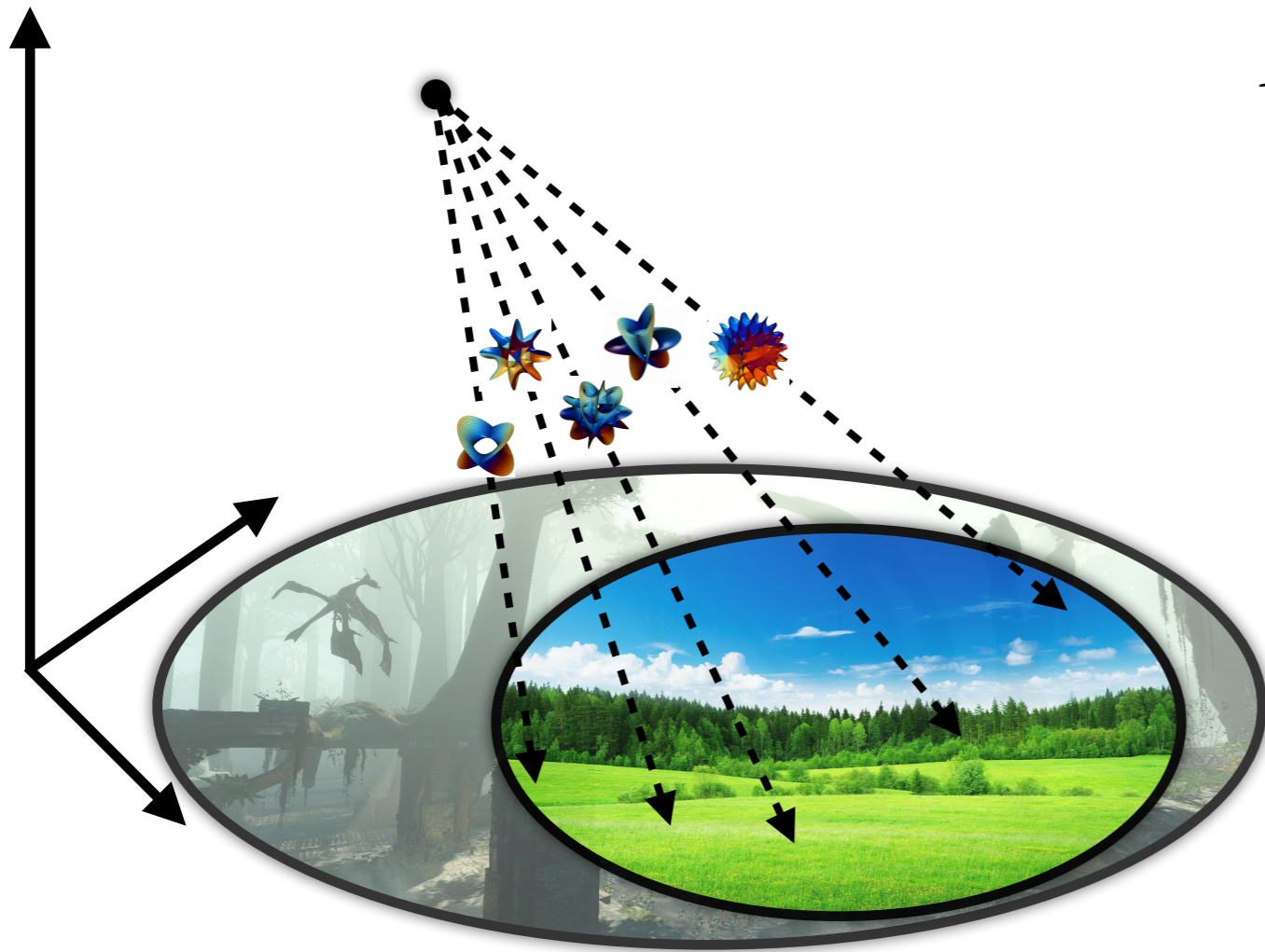


SWAMPLAND

LANDSCAPE



PART III: Theory



$$\mathcal{S} = \int d^4x \sqrt{-g} \left(\frac{R}{16\pi G_N} + \mathcal{L}_{\text{SM}} - \frac{1}{4} F_{\mu\nu} F^{\mu\nu} + \dots \right)$$

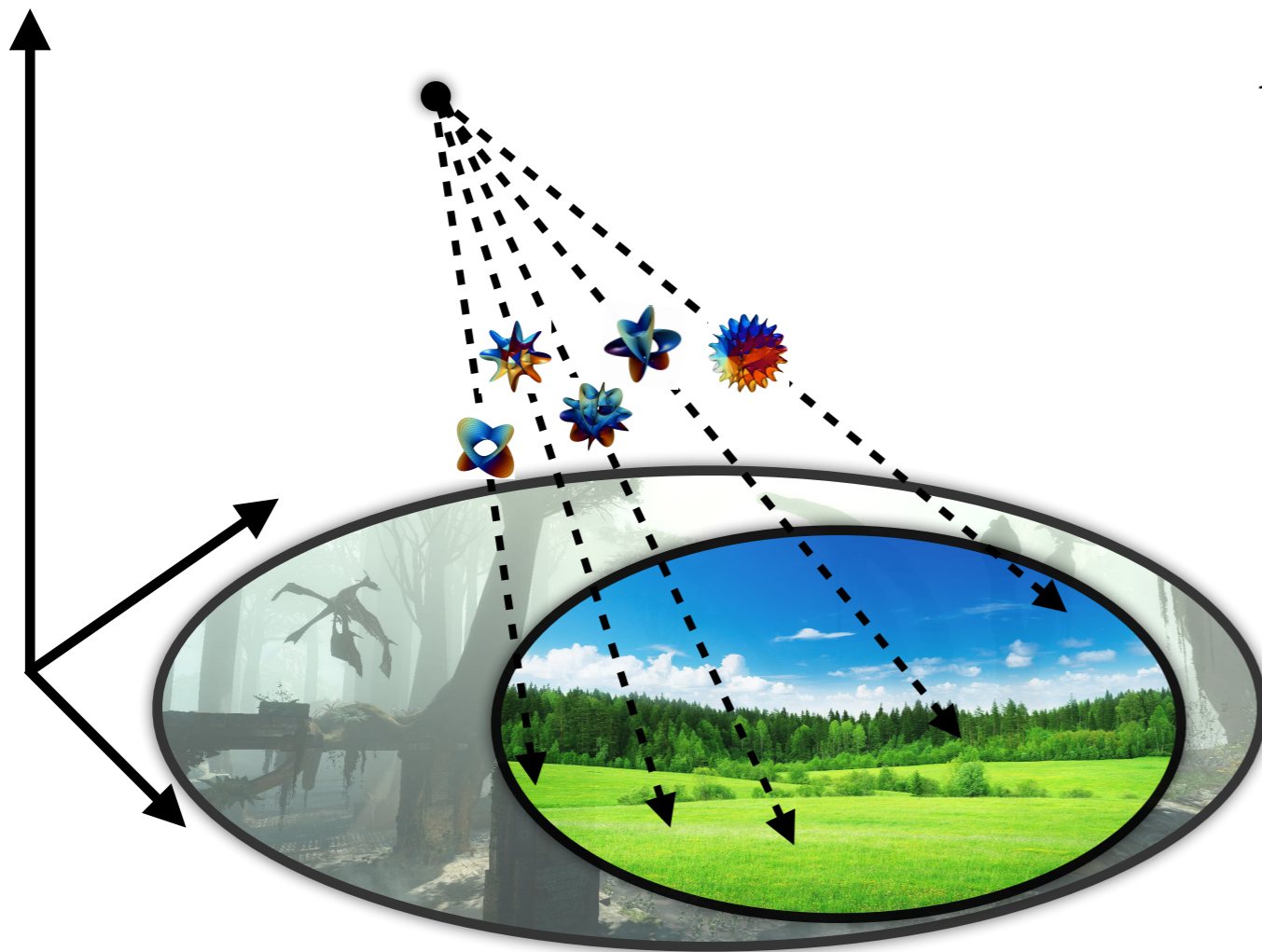
$U(1)$ dark

PART III: Theory

Weak Gravity Conjecture

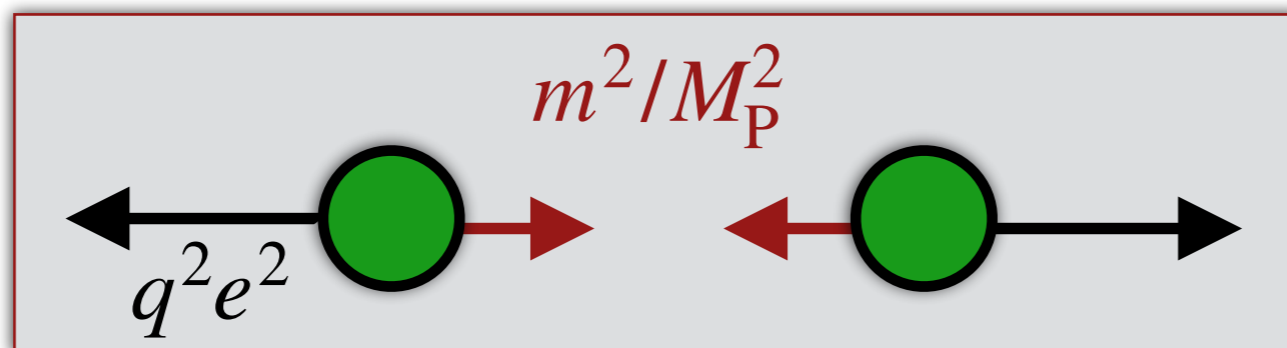
There must exist at least one state with charge qe larger than its mass m in Planck units

$$qe > \frac{m}{M_{\text{P}}}$$



$$\mathcal{S} = \int d^4x \sqrt{-g} \left(\frac{R}{16\pi G_N} + \mathcal{L}_{\text{SM}} - \frac{1}{4} F_{\mu\nu} F^{\mu\nu} + \dots \right)$$

$U(1)$ dark

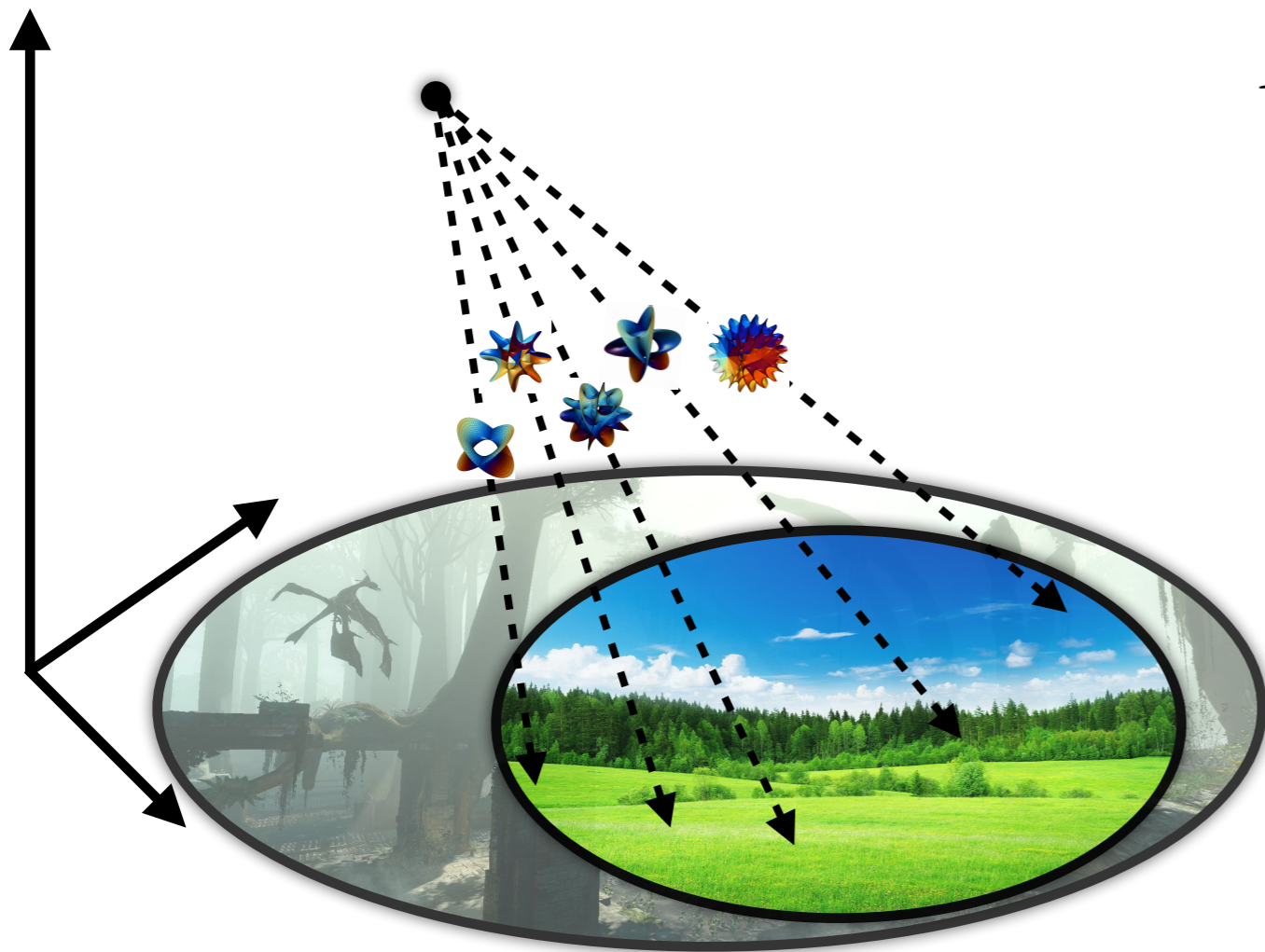


PART III: Theory

Weak Gravity Conjecture

There must exist at least one state with charge qe larger than its mass m in Planck units

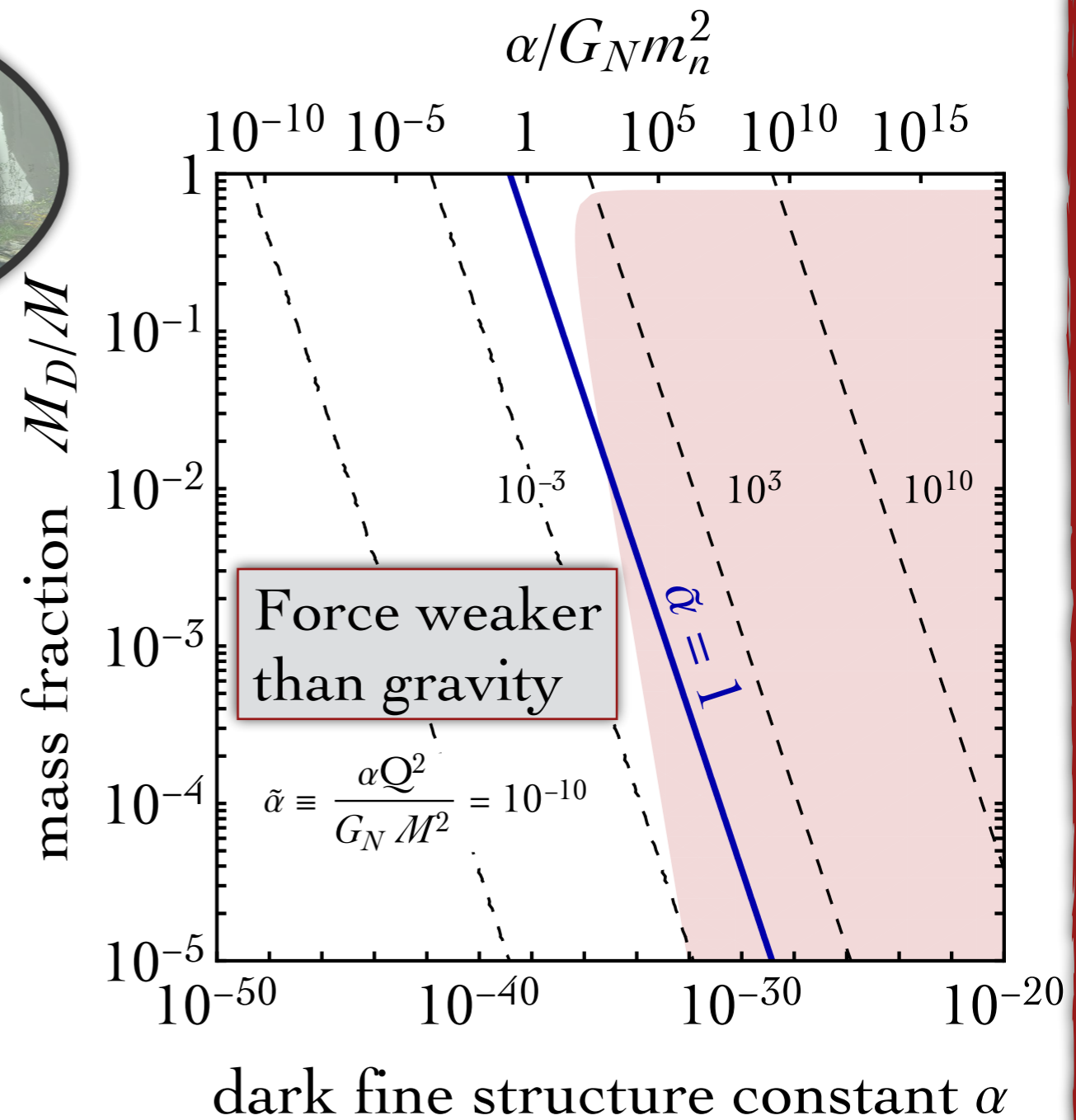
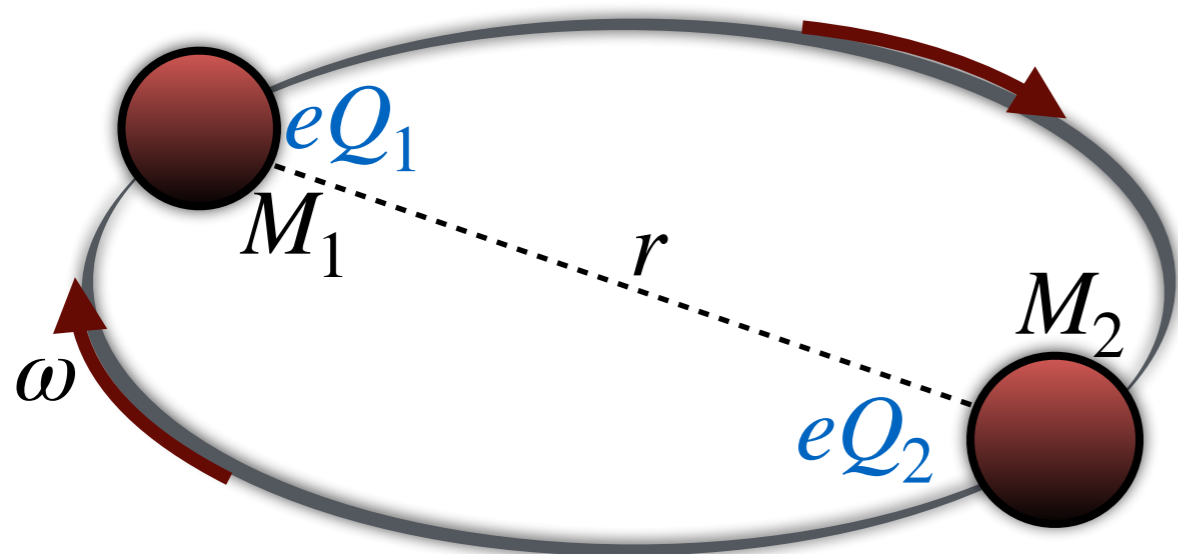
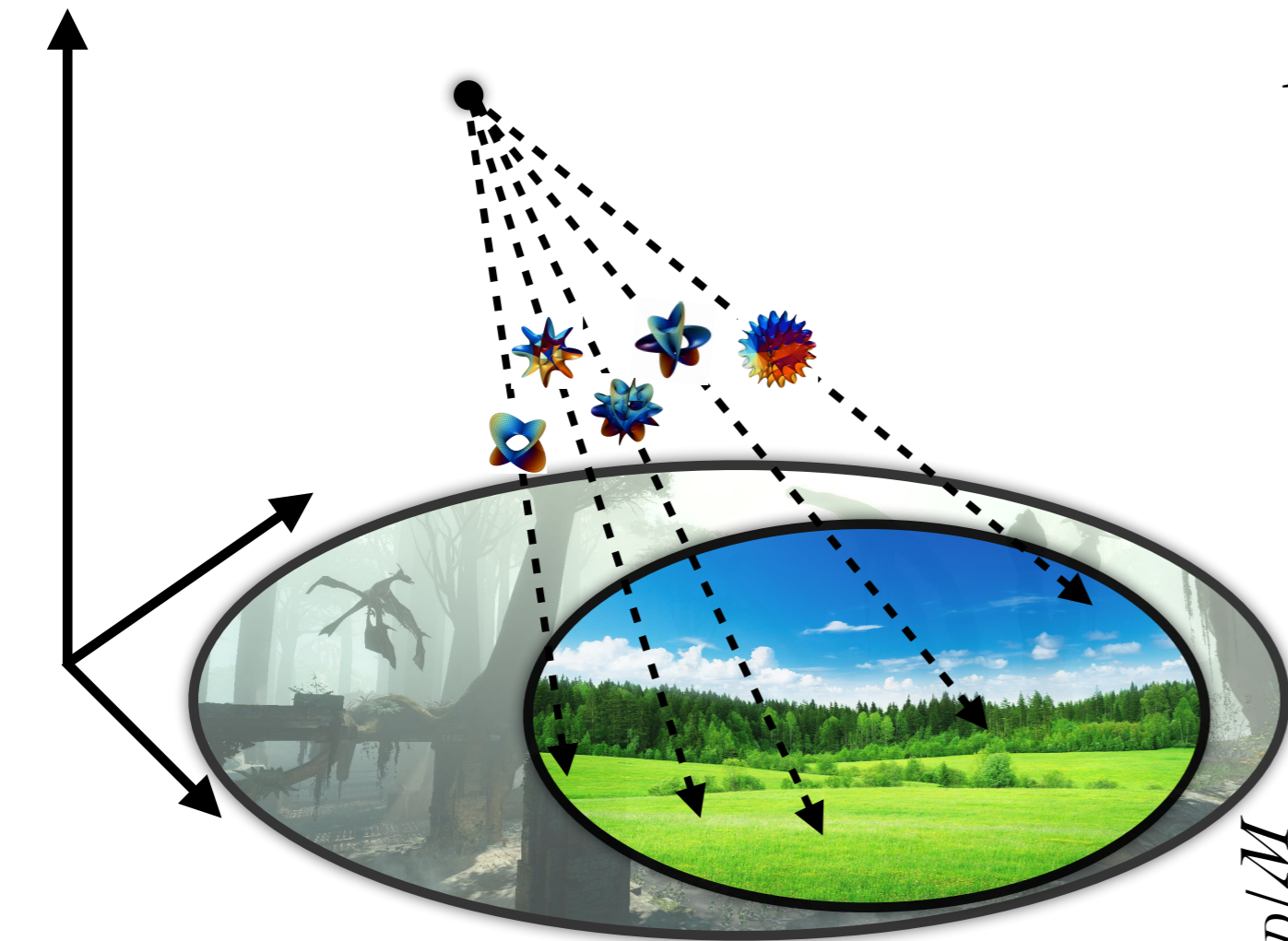
$$qe > \frac{m}{M_P}$$



It is tempting to generalize it...

“Gravity is the weakest force in Nature.”

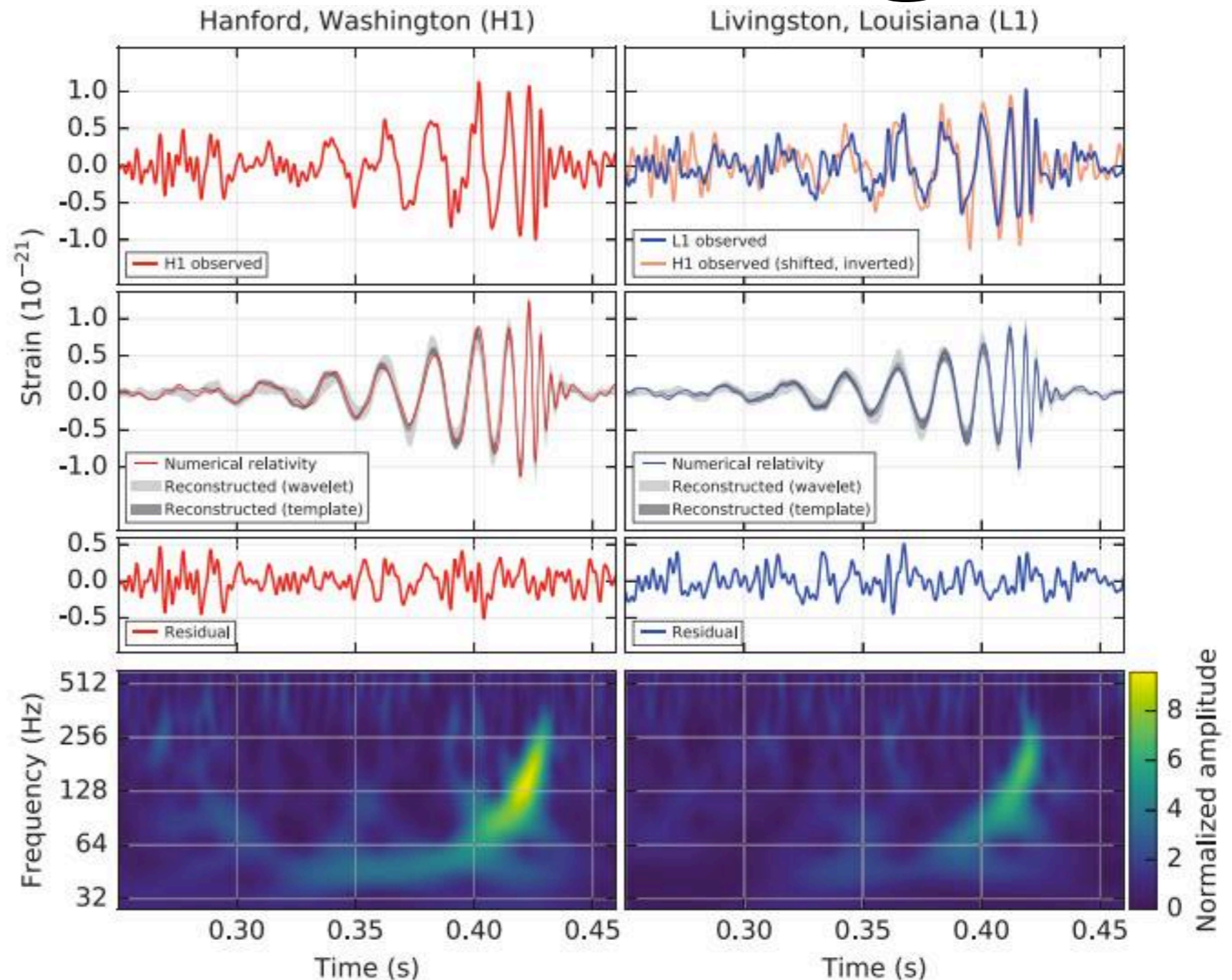
PART III: Theory



PART IV: Conclusions

PART IV: Conclusions

A new era has begun



ADDENDUM: References

Black hole superradiance: R. Brito, V. Cardoso and P. Pani, Lect. Notes Phys. 906, pp.1 (2015) [arXiv:1501.06570]; A. Arvanitaki and S. Dubovsky, Phys. Rev. D **83**, 044026 (2011) [arXiv:1004.3558]; A. D. Plascencia and AU, JCAP **1804**, no. 04, 059 (2018) [arXiv:1711.08298].

GW “echoes”: V. Cardoso, S. Hopper, C. F. B. Macedo, C. Palenzuela and P. Pani, Phys. Rev. D **94**, no. 8, 084031 (2016) [arXiv:1608.08637]; V. Cardoso and P. Pani, Nat. Astron. **1**, no. 9, 586 (2017) [arXiv:1709.01525]; P. Bueno *et al.*, Phys. Rev. D **97**, no. 2, 024040 (2018) [arXiv:1711.00391]; P. Pani and V. Ferrari, Class. Quant. Grav. **35**, no. 15, 15LT01 (2018) [arXiv:1804.01444]; AU and H. Veermäe, arXiv:1810.07137.

Neutron star EoS from GW: E. Annala, T. Gorda, A. Kurkela and A. Vuorinen, Phys. Rev. Lett. **120** (2018) no.17, 172703 [arXiv:1711.02644]; E. R. Most, L. R. Weih, L. Rezzolla and J. Schaffner-Bielich, Phys. Rev. Lett. **120**, no. 26, 261103 (2018) [arXiv:1803.00549].

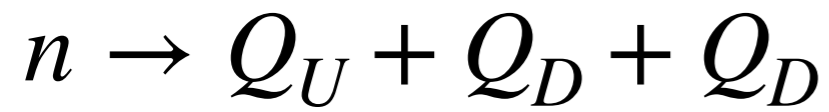
Post-Newtonian expansion and scattering amplitudes: R. A. Porto, Phys. Rev. D **73**, 104031 (2006) [gr-qc/0511061]; C. Cheung, I. Z. Rothstein and M. P. Solon, Phys. Rev. Lett. **121**, no. 25, 251101 (2018) [arXiv:1808.02489].

Charged neutron stars and GW: L. Sagunski *et al.*, Phys. Rev. D **97**, no. 6, 064016 (2018) [arXiv:1709.06634]; D. Croon *et al.*, Astrophys. J. **858**, no. 1 (2018), L2 [arXiv:1711.02096]; J.~Kopp *et al.*, JHEP **1811**, 096 (2018) [arXiv:1807.02527]; M. Fabbrichesi and AU (to appear).

The swampland conjecture: C. Vafa, hep-th/0509212; H. Ooguri and C. Vafa, Adv. Theor. Math. Phys. **21**, 1787 (2017) [arXiv:1610.01533].

The weak gravity conjecture: N. Arkani-Hamed, L. Motl, A. Nicolis and C. Vafa, JHEP **0706**, 060 (2007) [hep-th/0601001]; B. Heidenreich, M. Reece and T. Rudelius, JHEP **1602**, 140 (2016) [arXiv:1509.06374]; S. Hod, Int. J. Mod. Phys. D **26**, no. 12, 1742004 (2017) [arXiv:1705.06287]; AU, arXiv:1810.05621; C. Cheung, J. Liu and G. N. Remmen, JHEP **1810**, 004 (2018) [arXiv:1801.08546]; M. Montero, arXiv:1812.03978.

Bonus Material



$$937.900 \text{ MeV} < 2m_D + m_U < m_n = 939.565 \text{ MeV}$$

decay of ${}^9\text{Be}$ closed

decay kin. open

